



# Dark Side of the Universe II

*Alexander Vikman*

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**FZU**

Institute of Physics  
of the Czech  
Academy of Sciences

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EUROPEAN UNION  
European Structural and Investment Funds  
Operational Programme Research,  
Development and Education

**MSMT**  
MINISTRY OF EDUCATION,  
YOUTH AND SPORTS

# Literature

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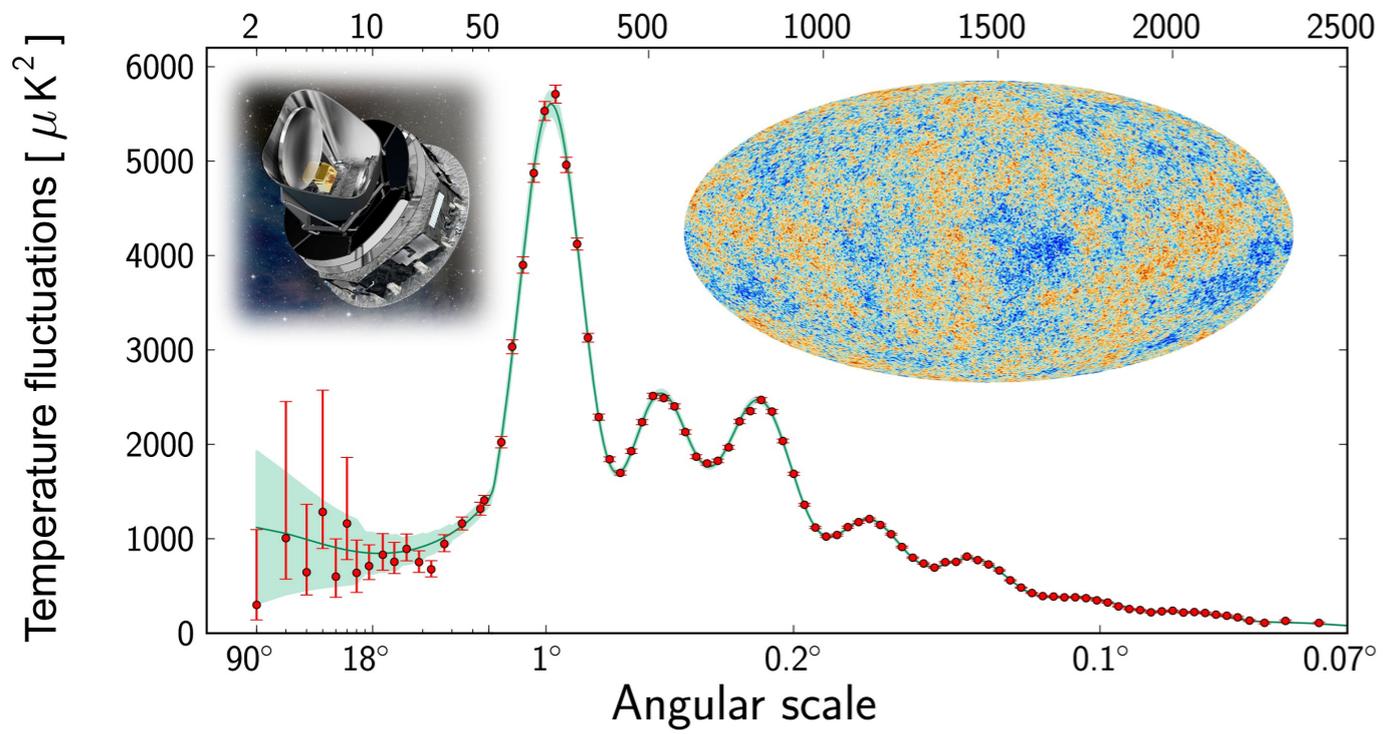
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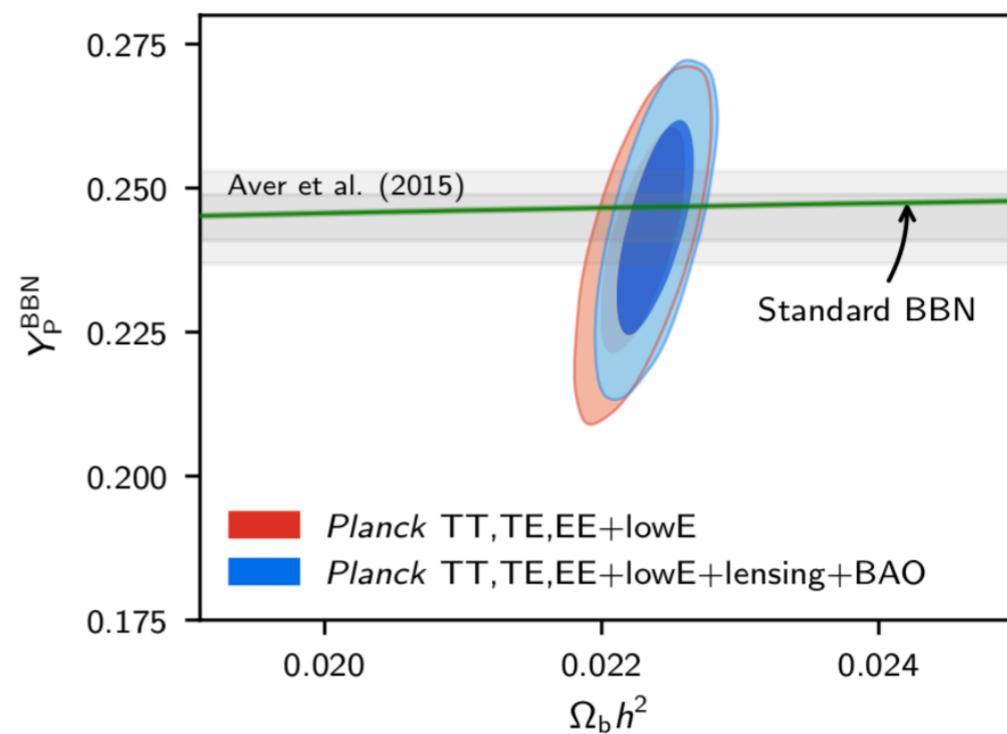
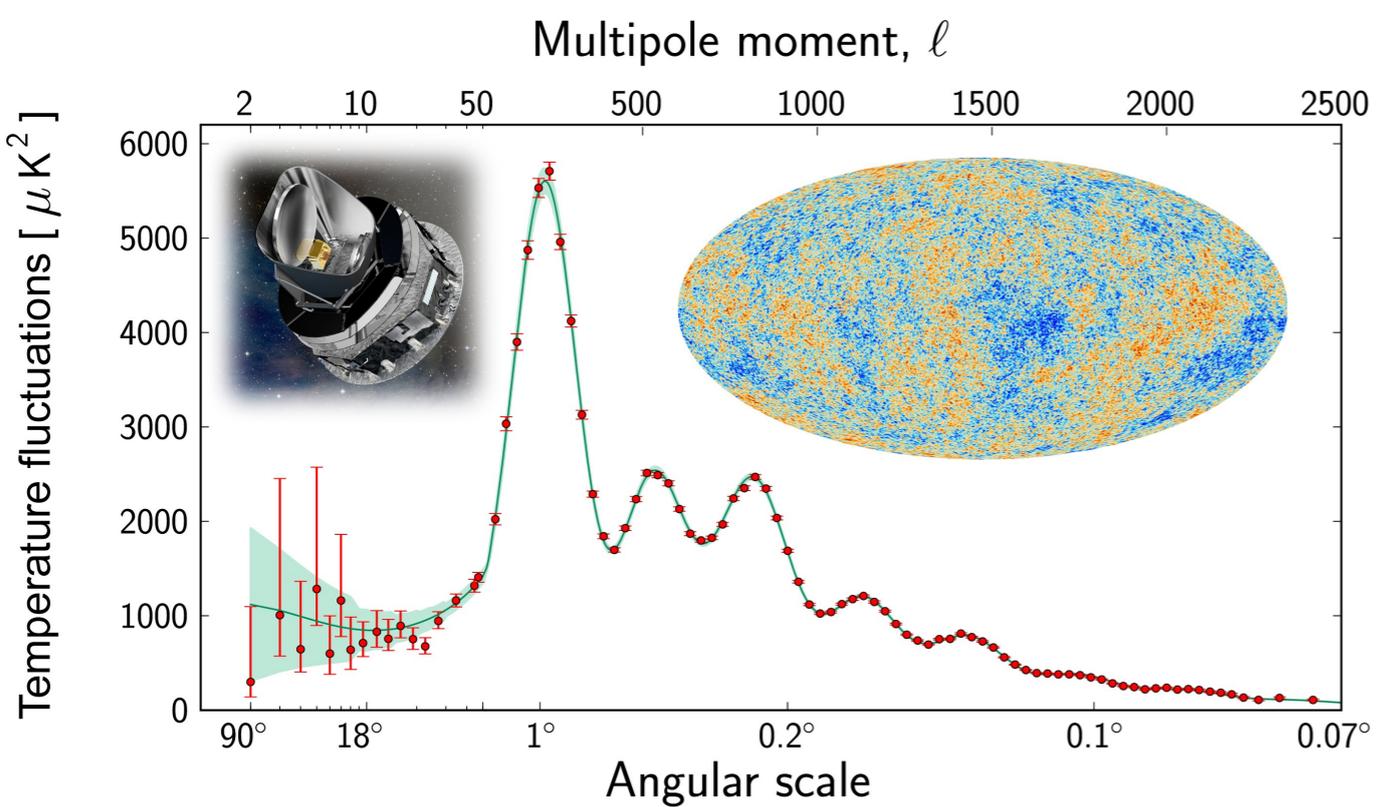
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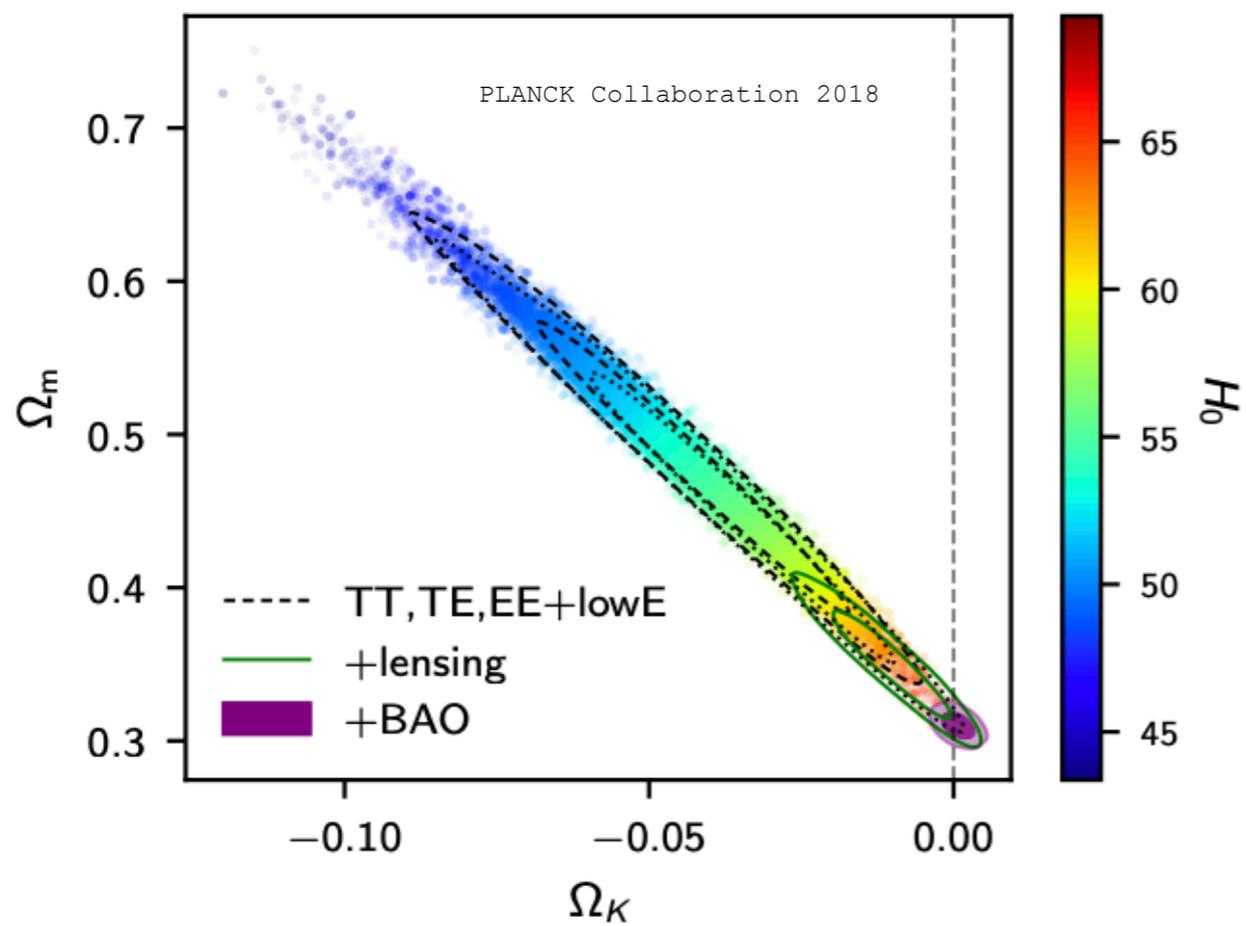
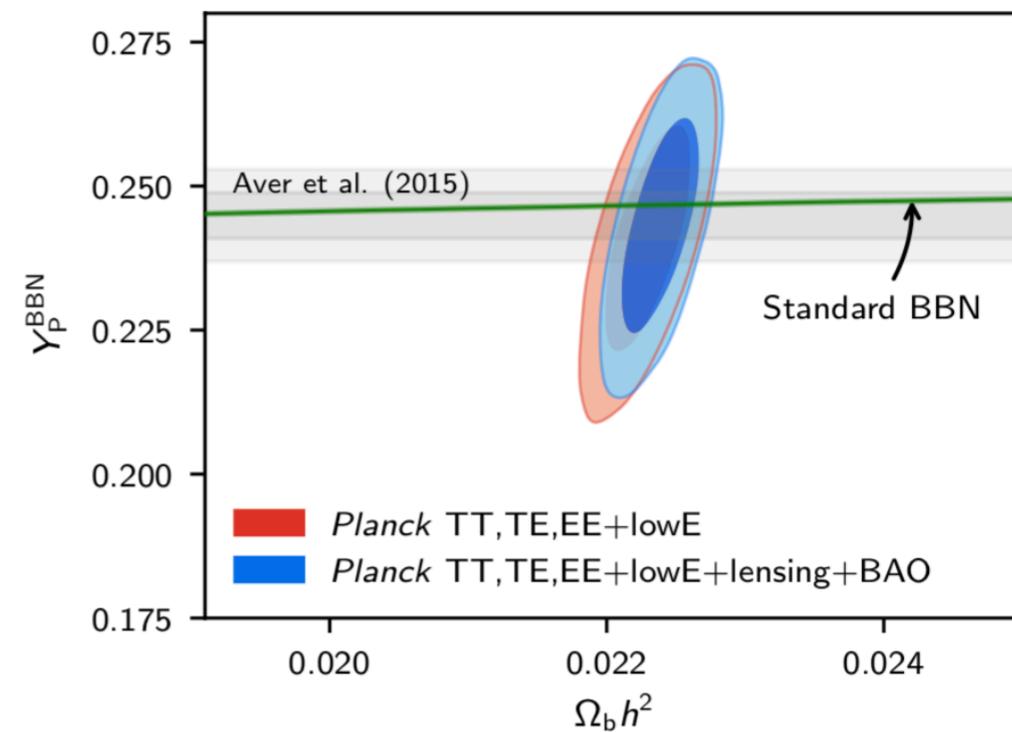
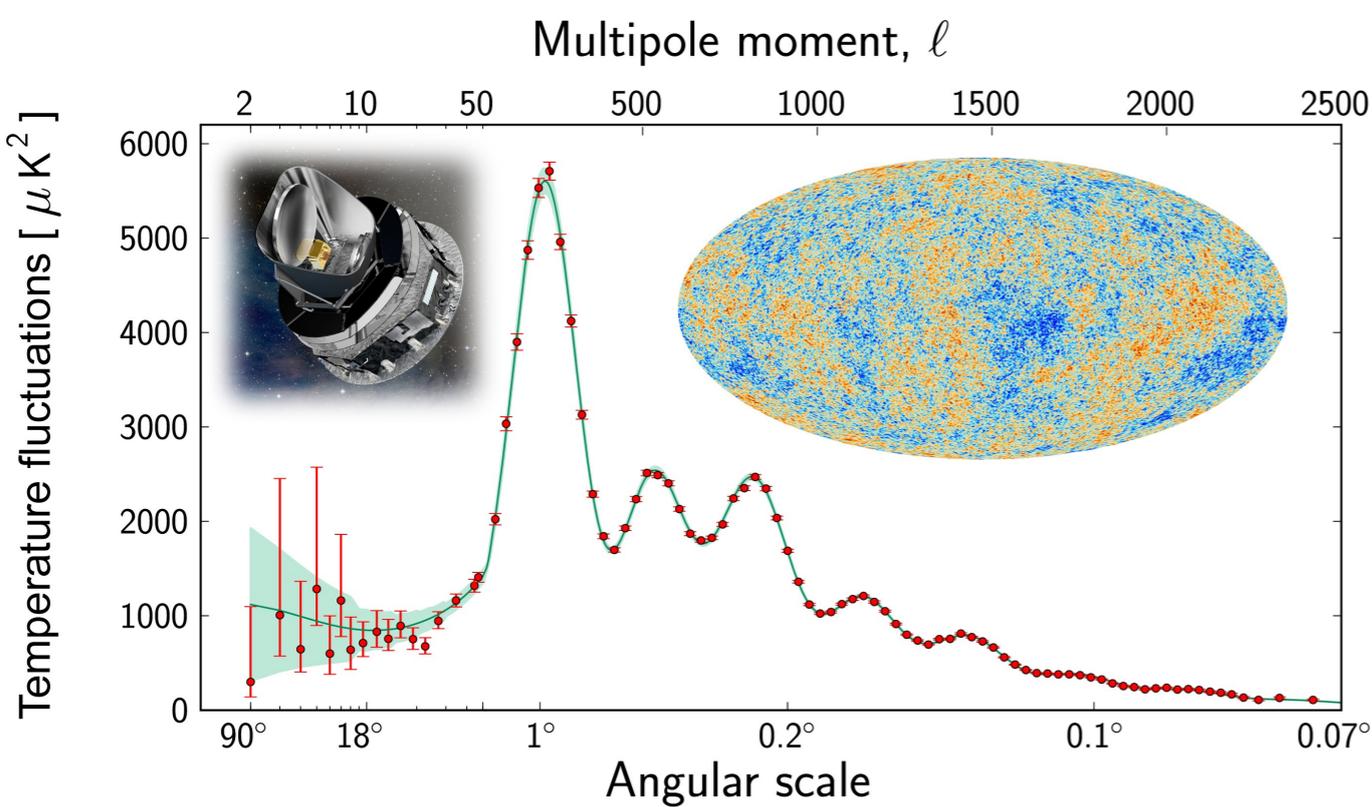
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- *Lectures on the Cosmological Constant Problem*  
A. Padilla  
arXiv:1502.05296

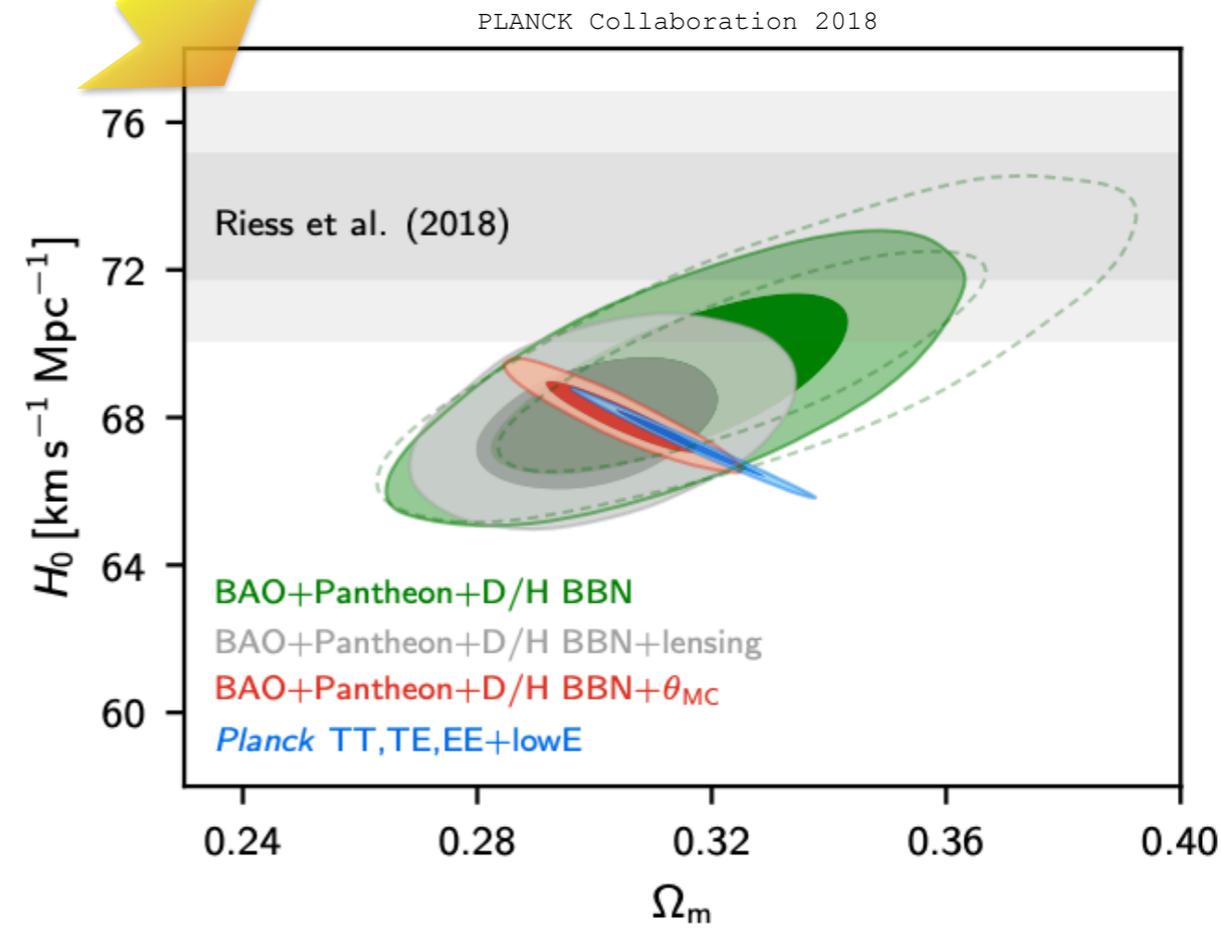
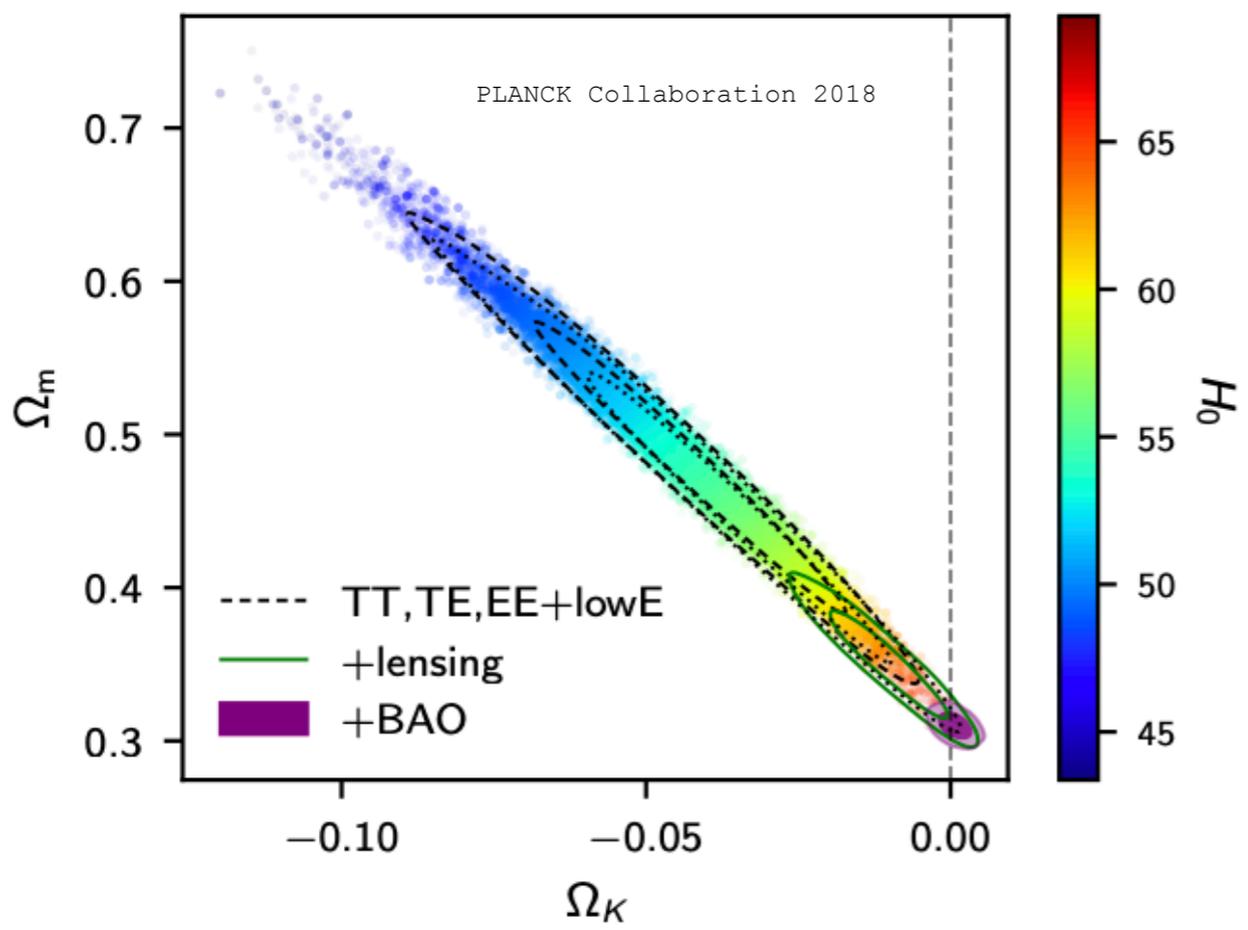
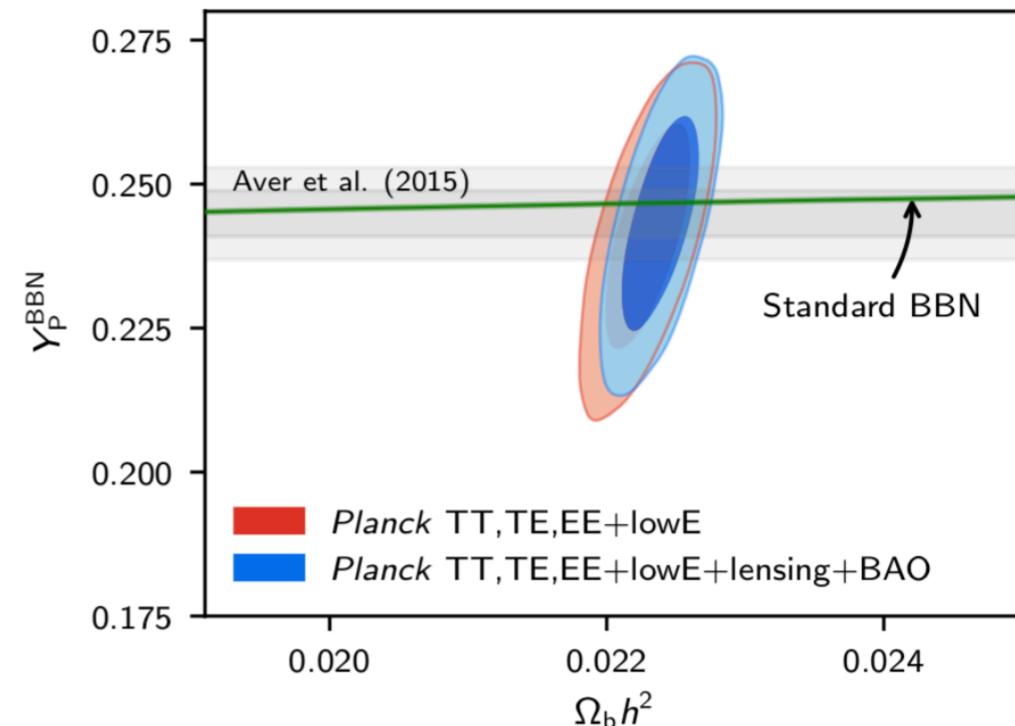
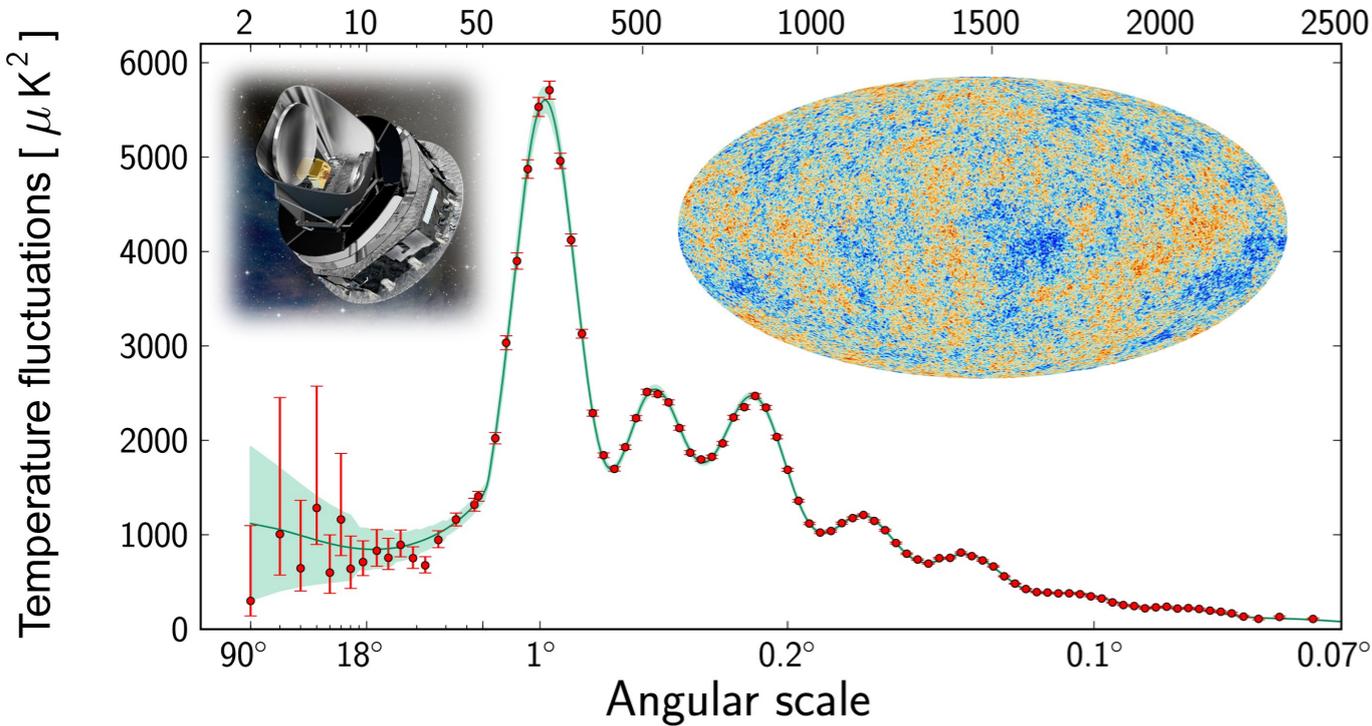
Multipole moment,  $l$

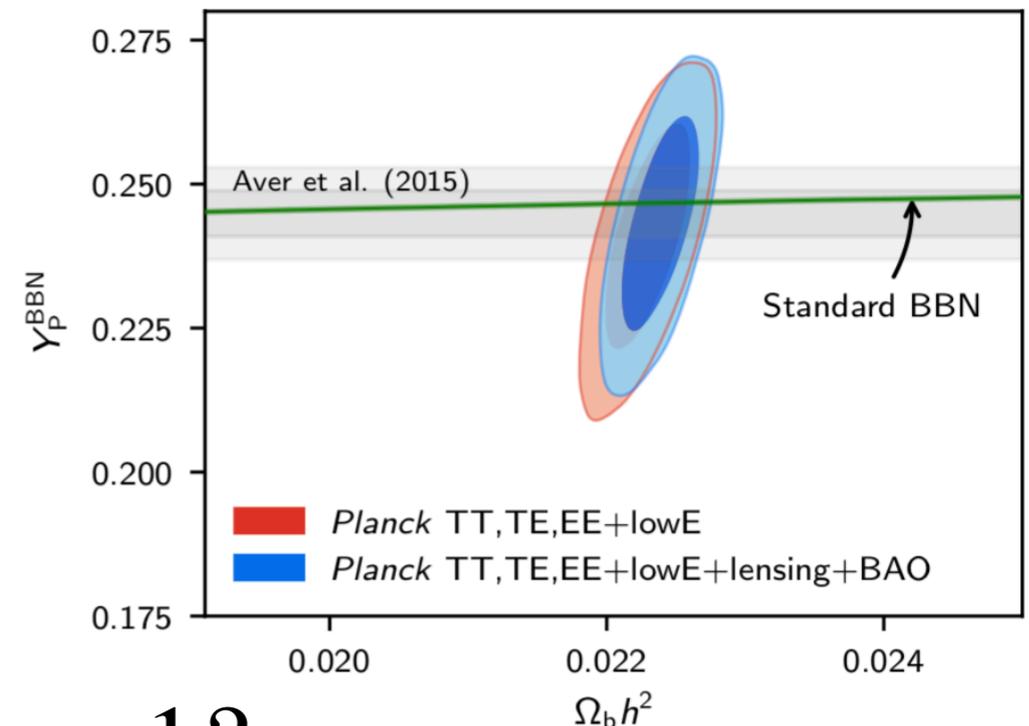
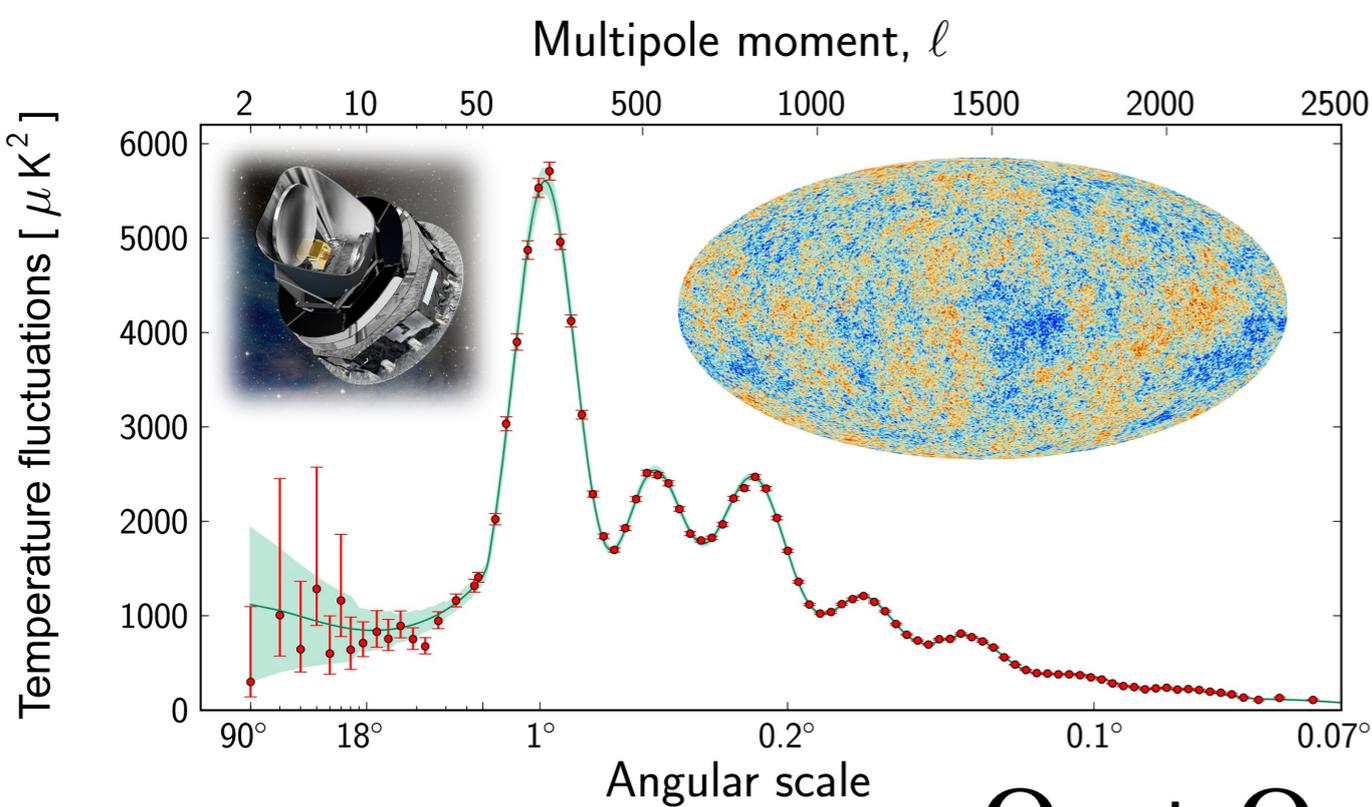




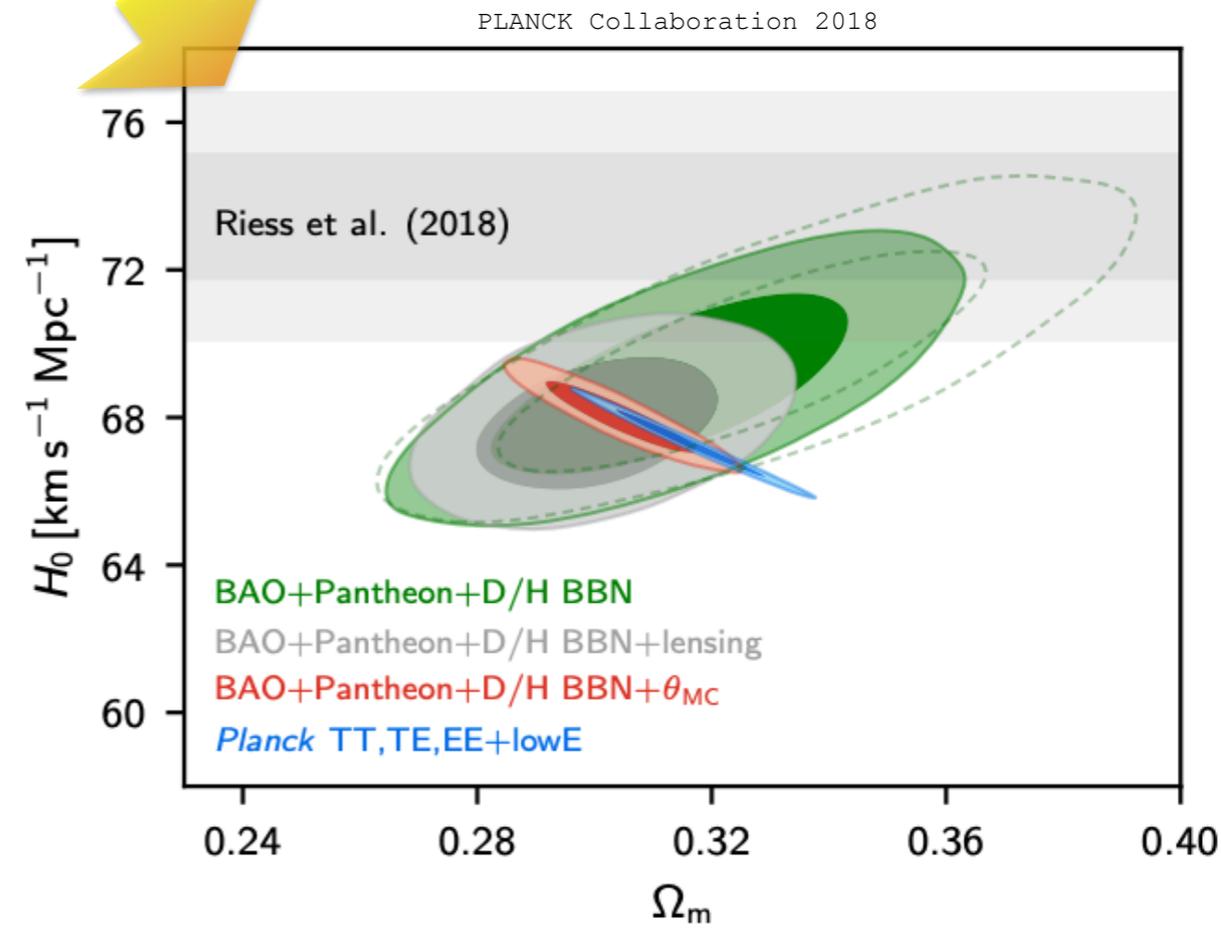
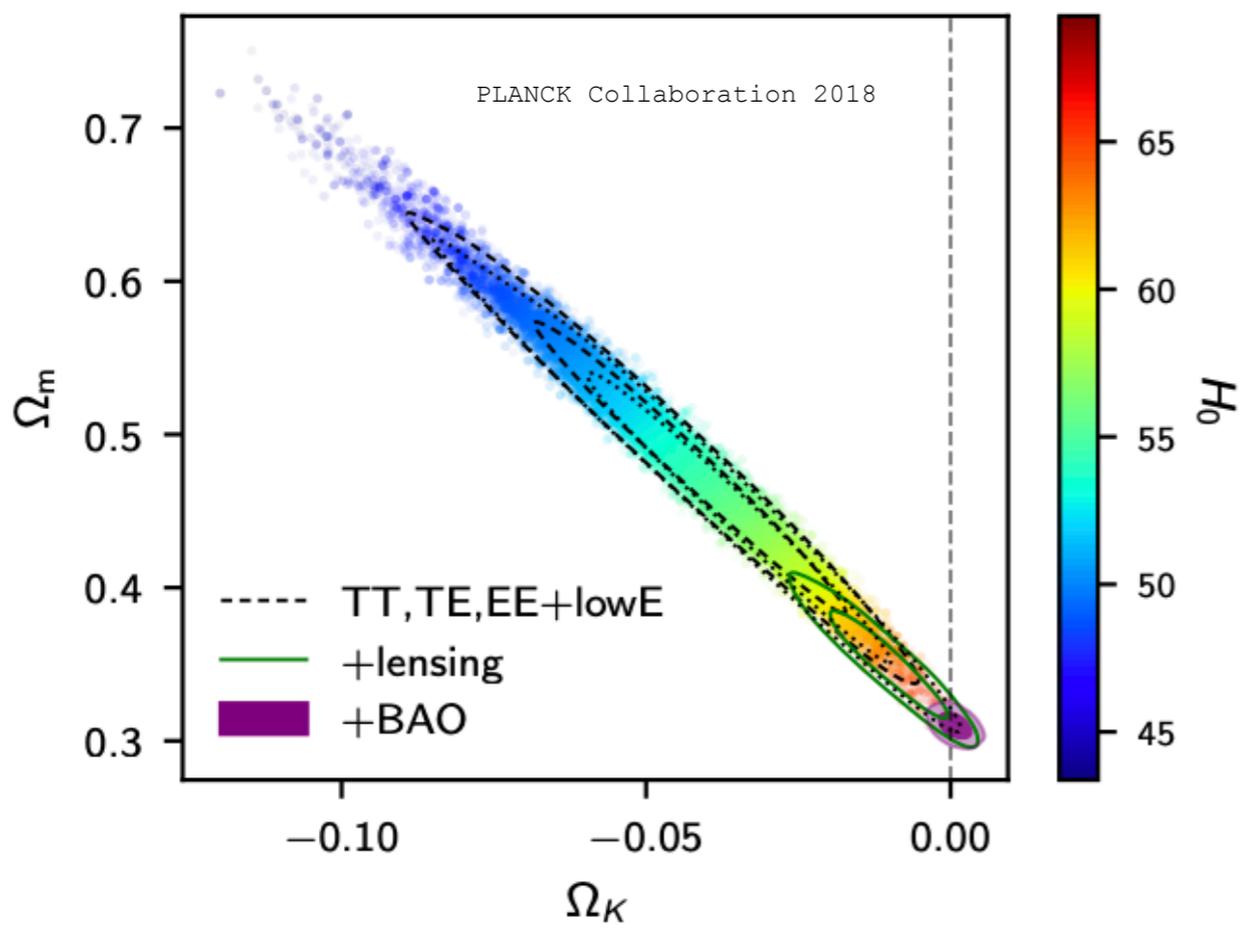


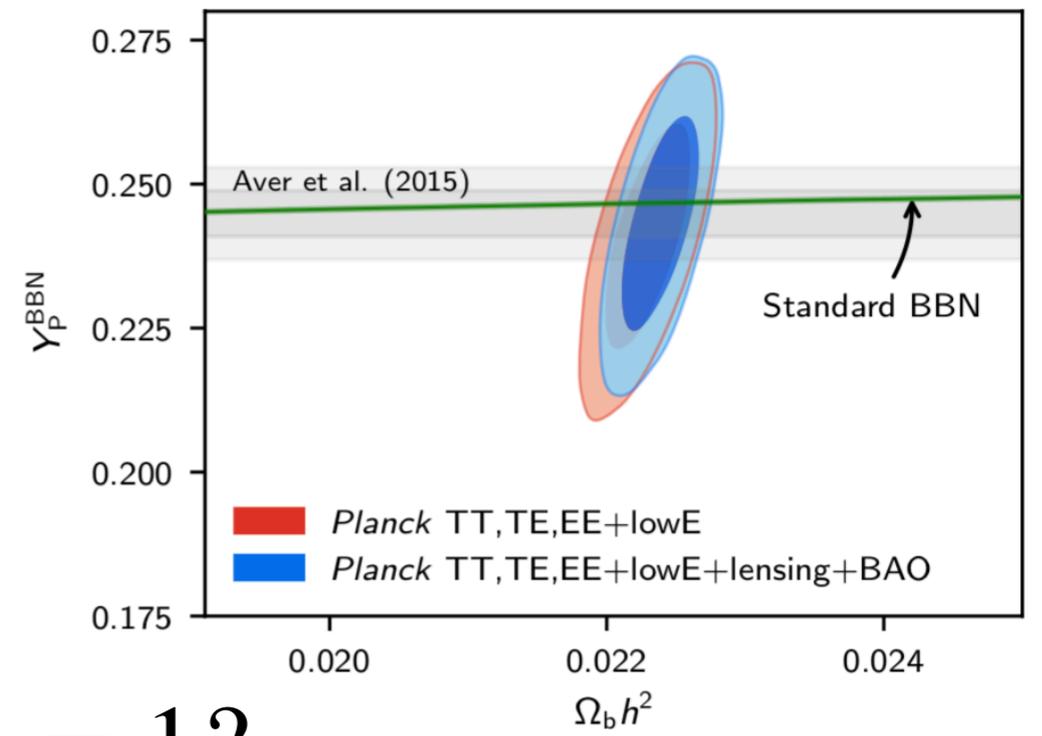
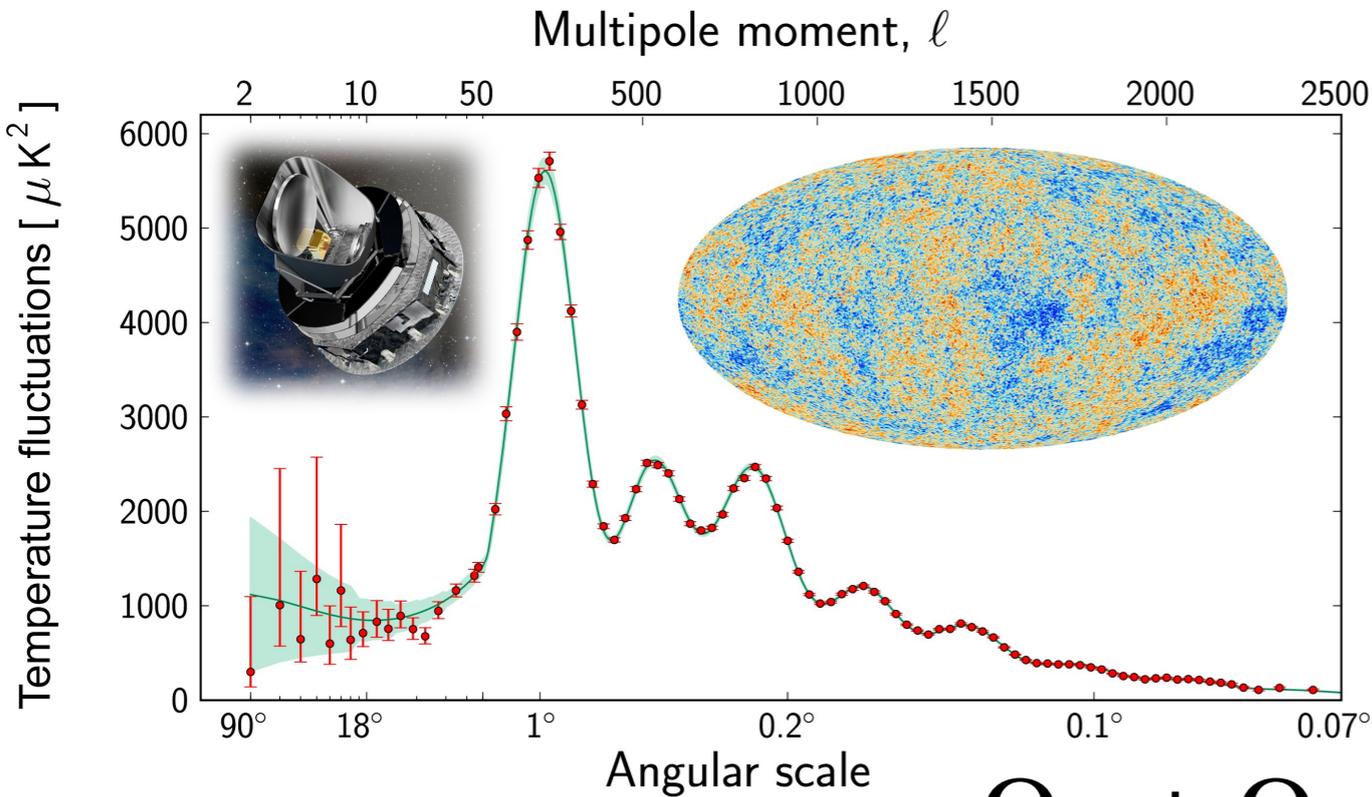
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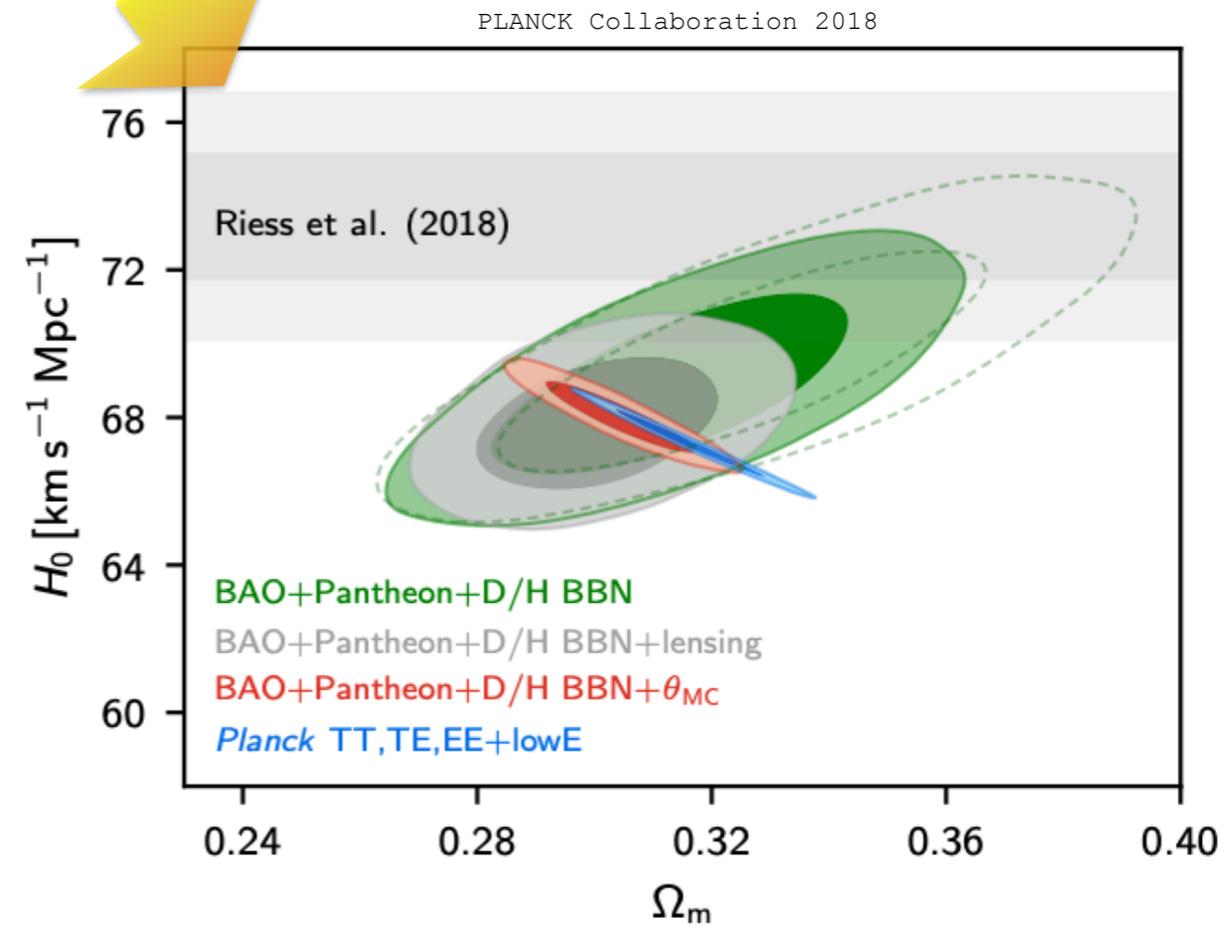
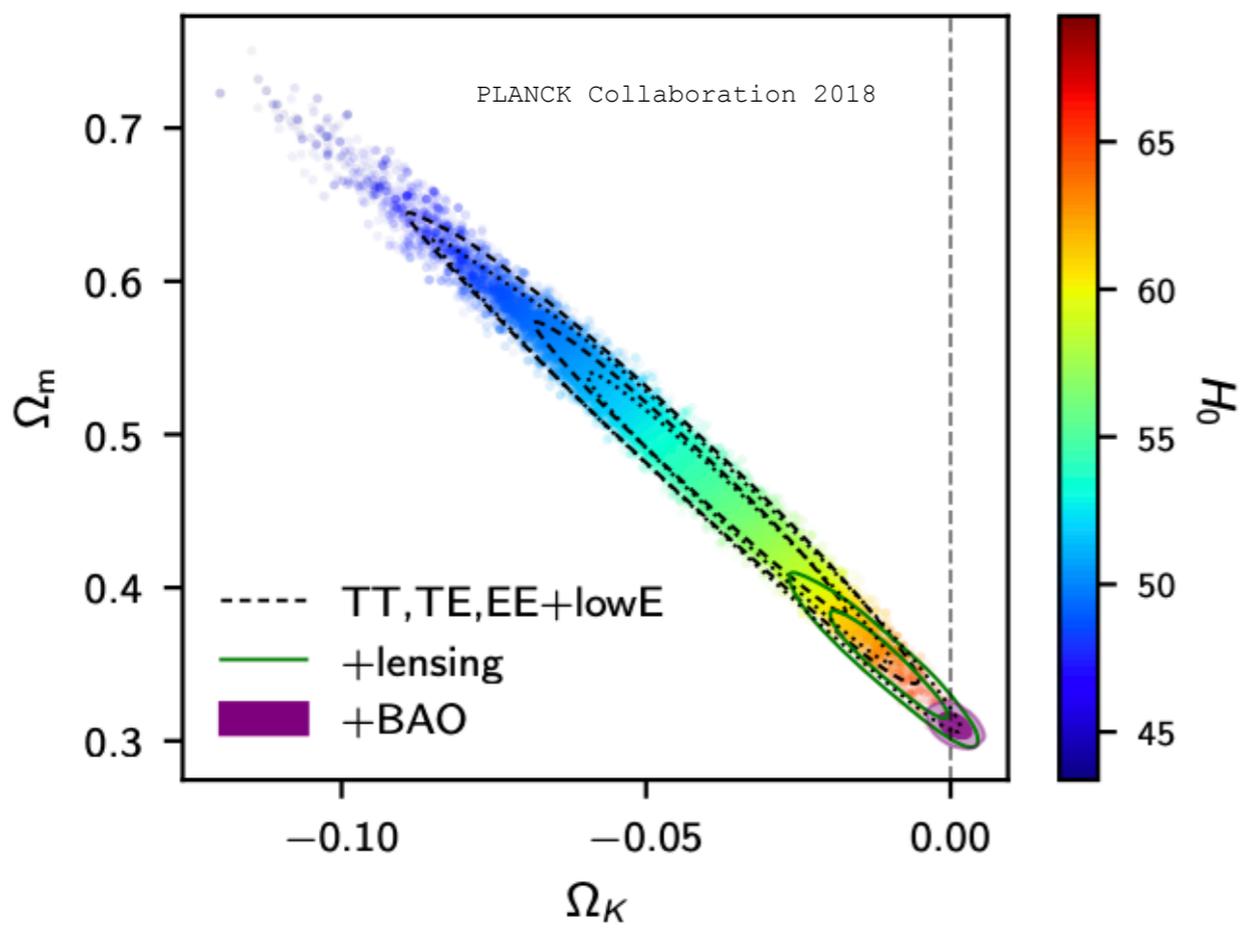
$$\Omega_m + \Omega_b + \Omega_k = 1?$$

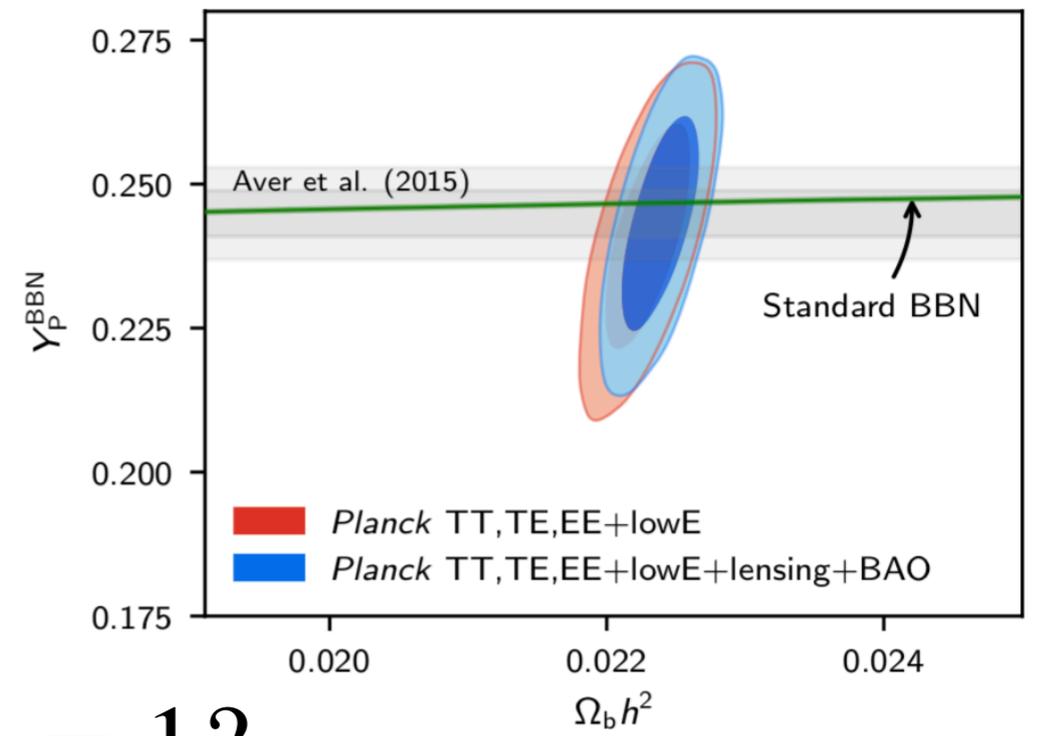
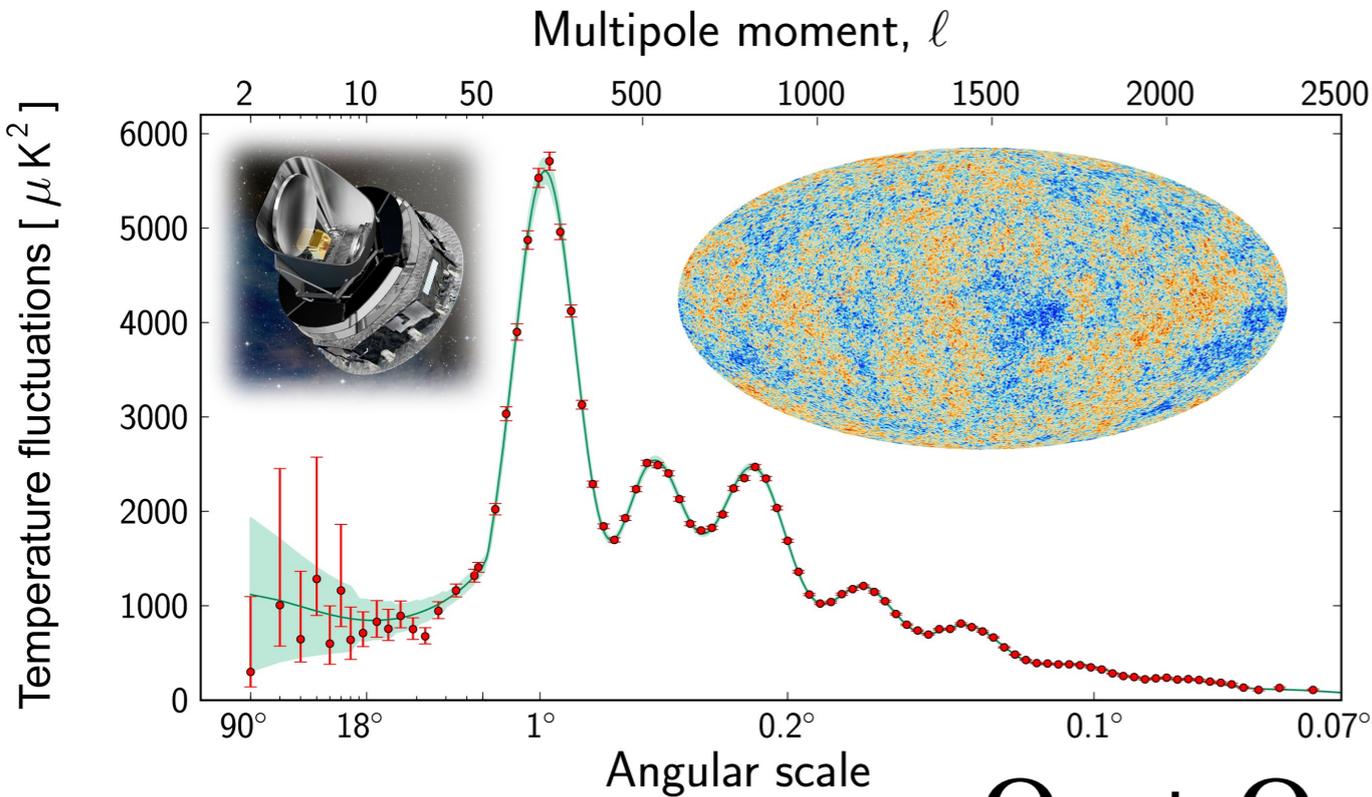




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$$\Omega_m + \Omega_b + \Omega_k \simeq 0.32$$

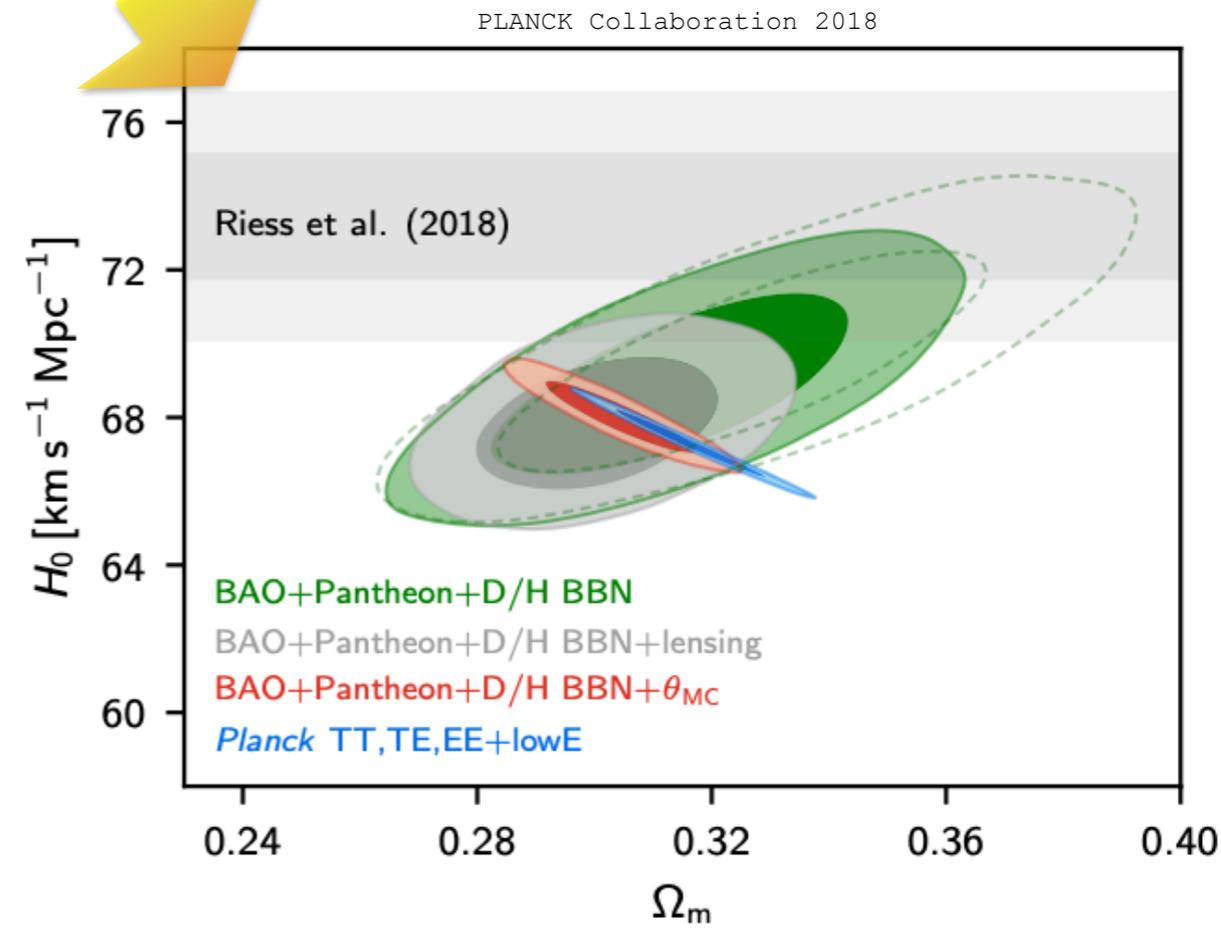
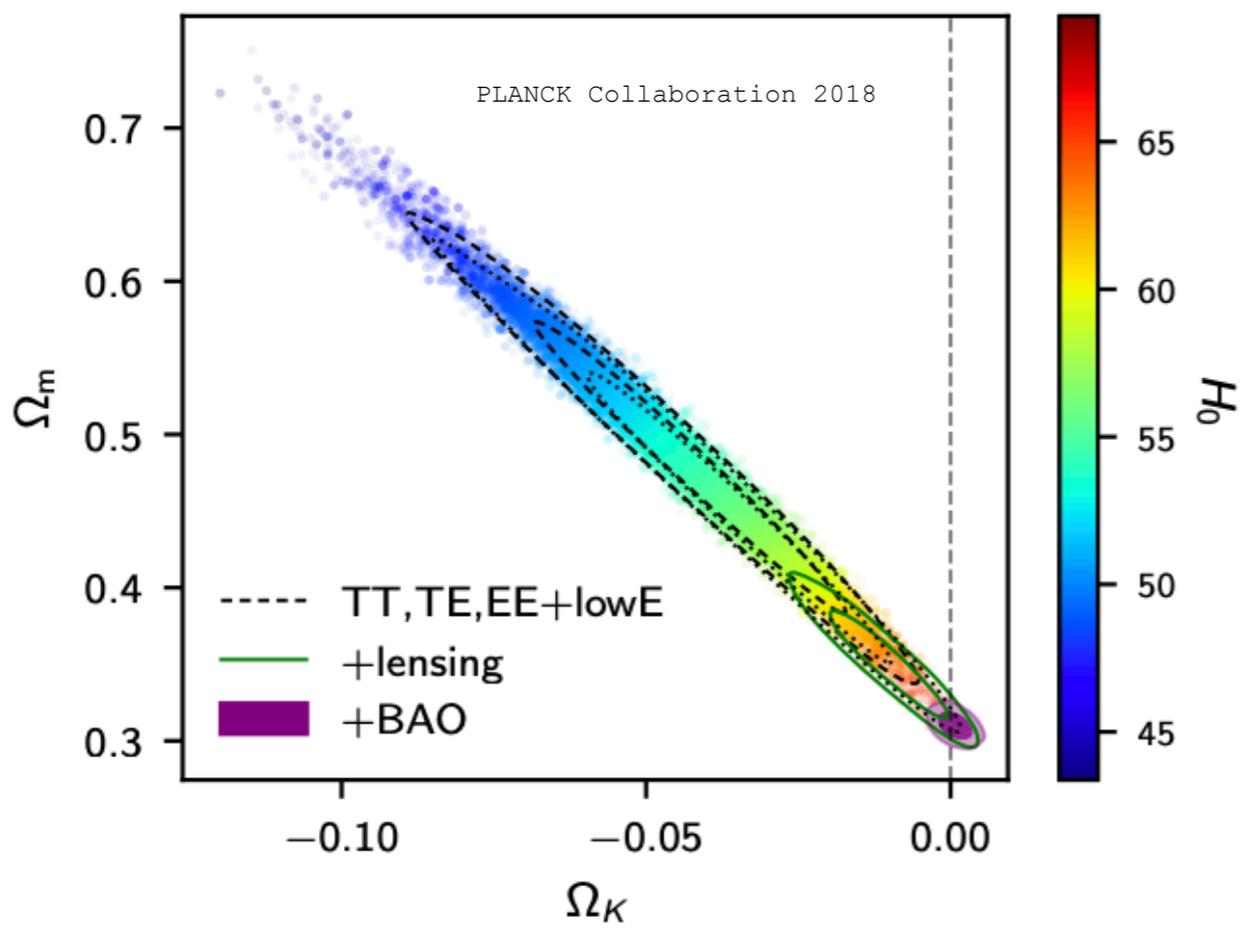




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**Dark Energy!**



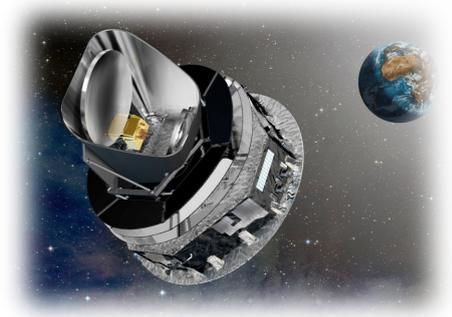
# Dark Energy equation of state

$$w = p/\varepsilon \quad w(a) = w_0 + (1 - a) w_a$$



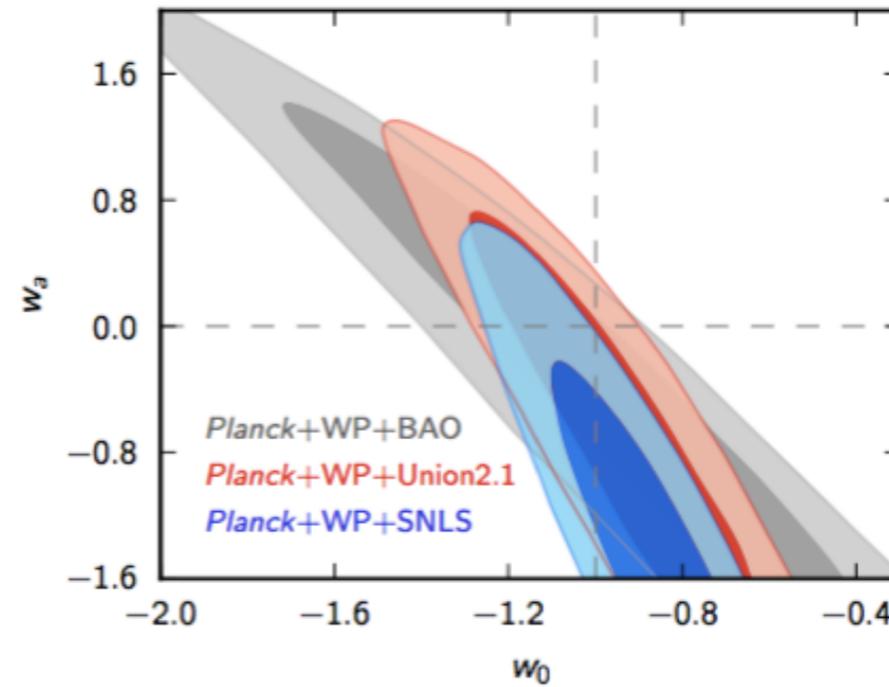
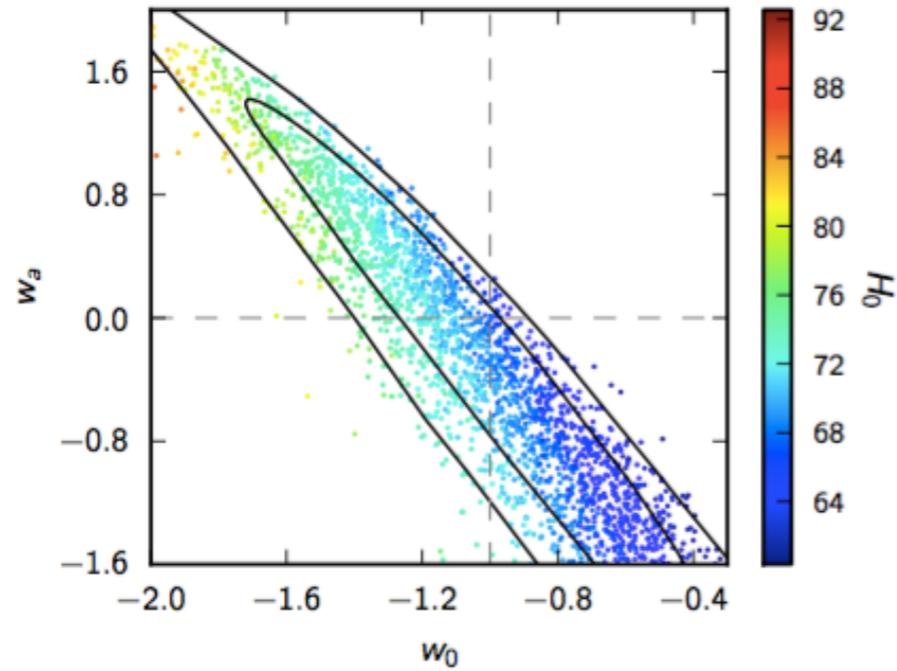
PLANCK Collaboration 2013

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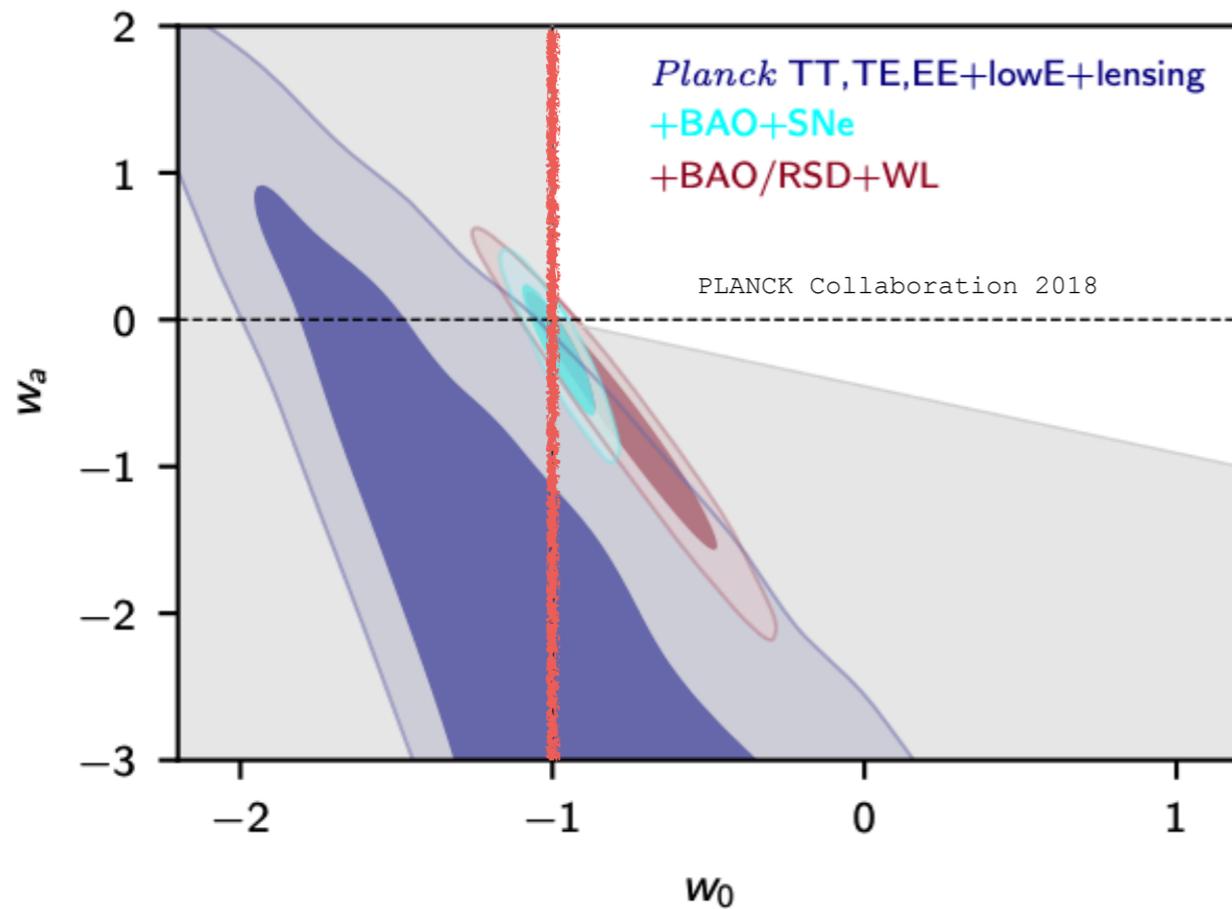
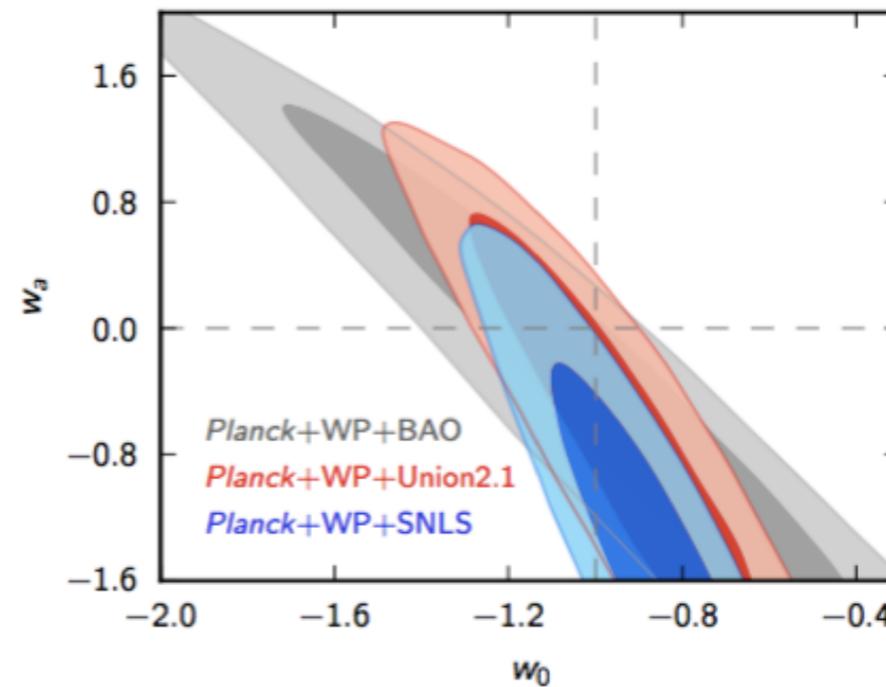
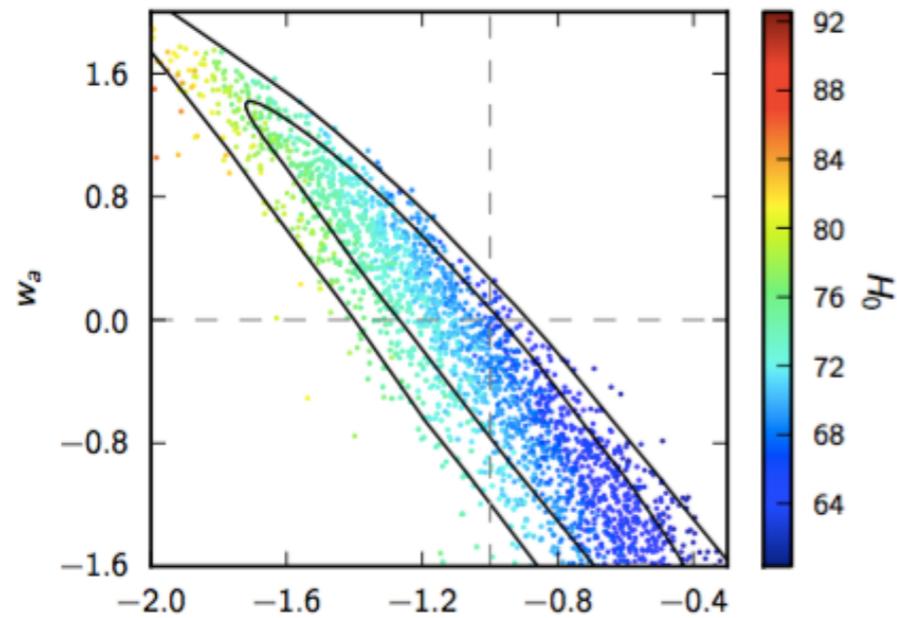


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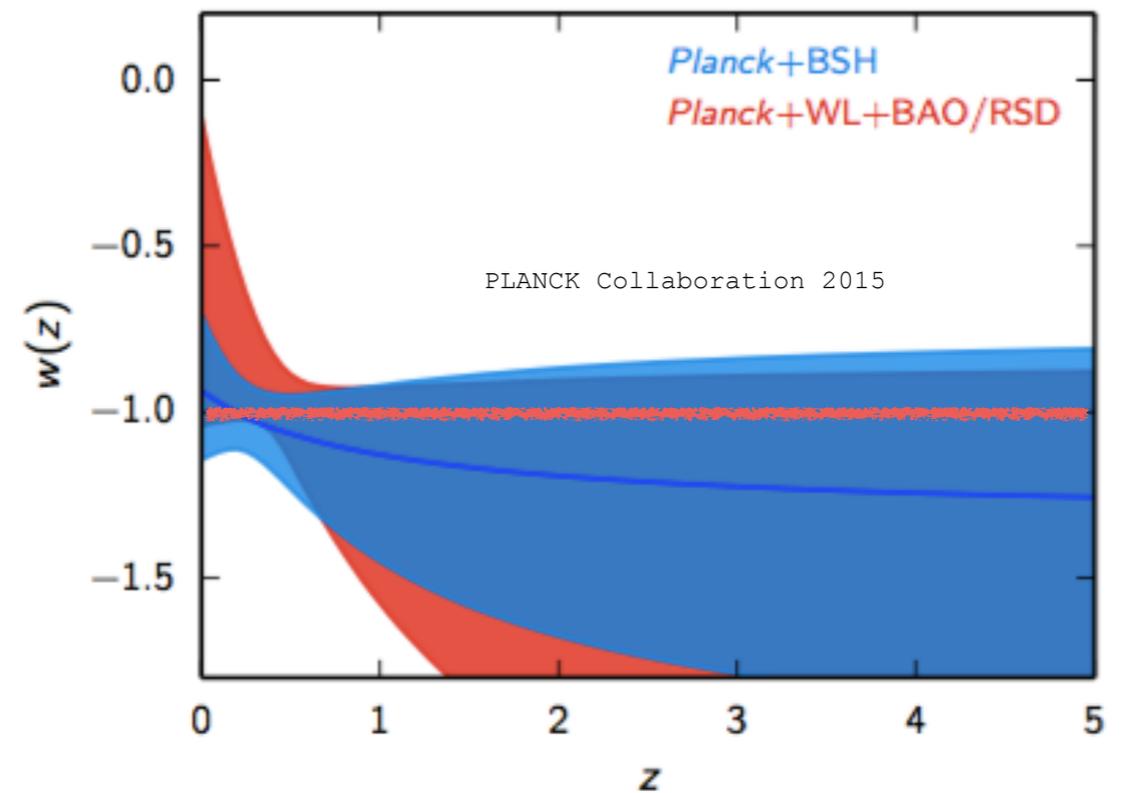
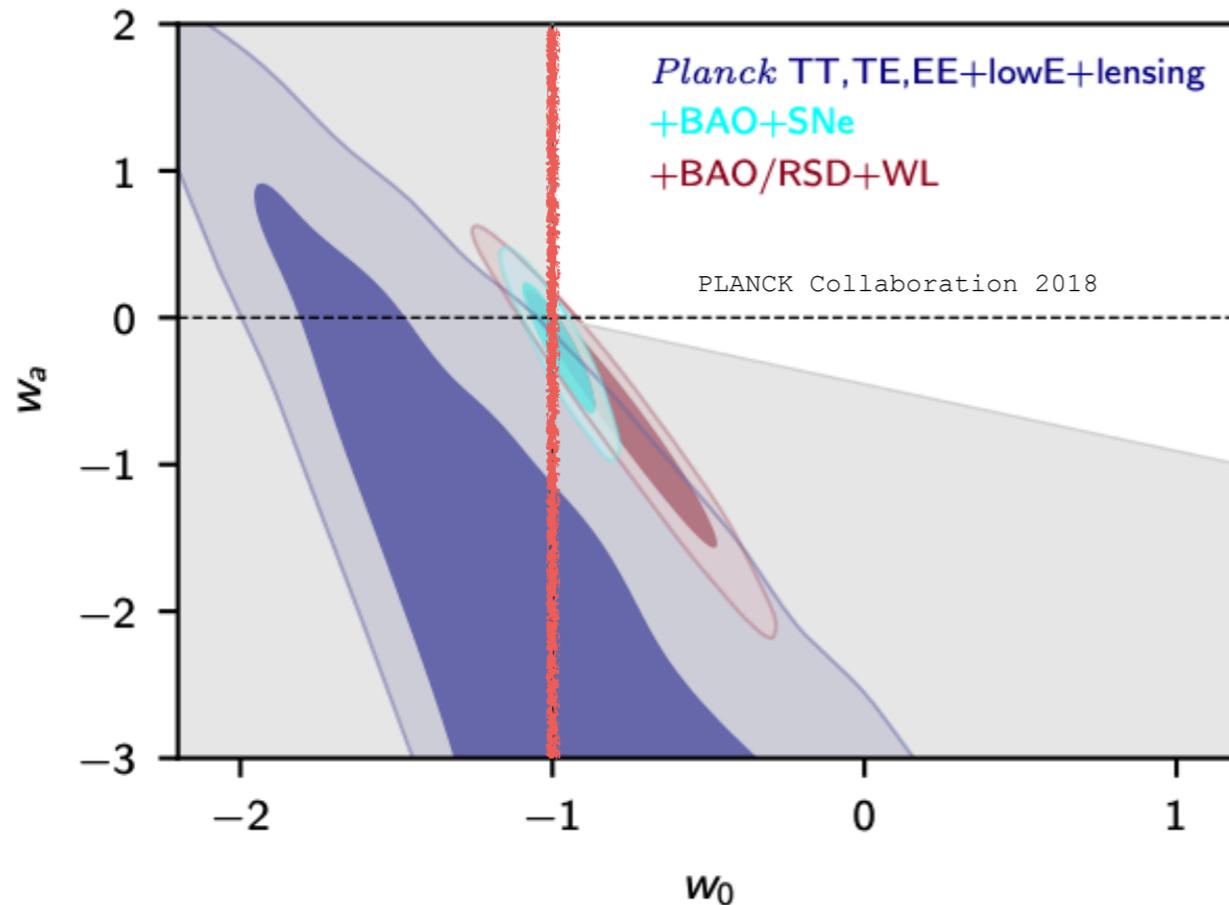
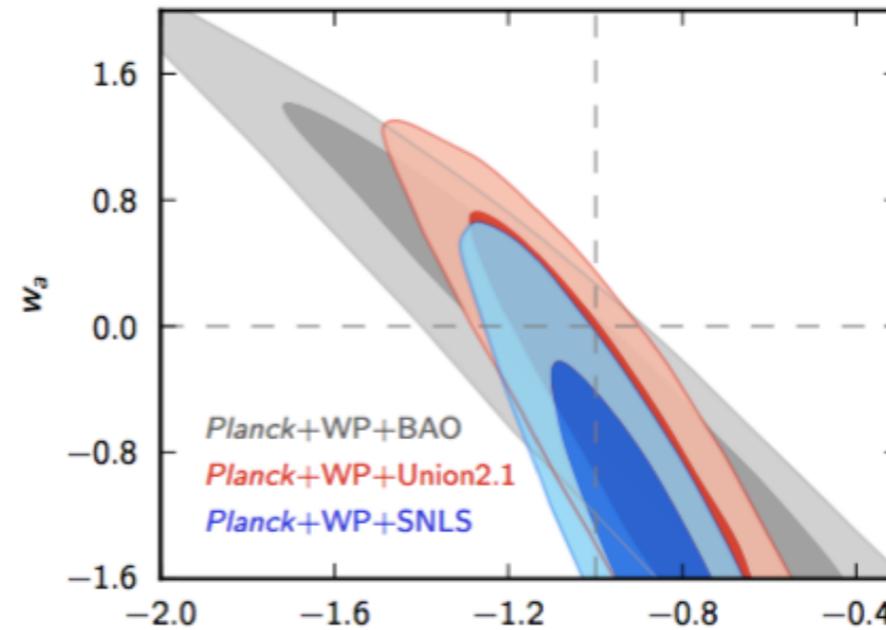
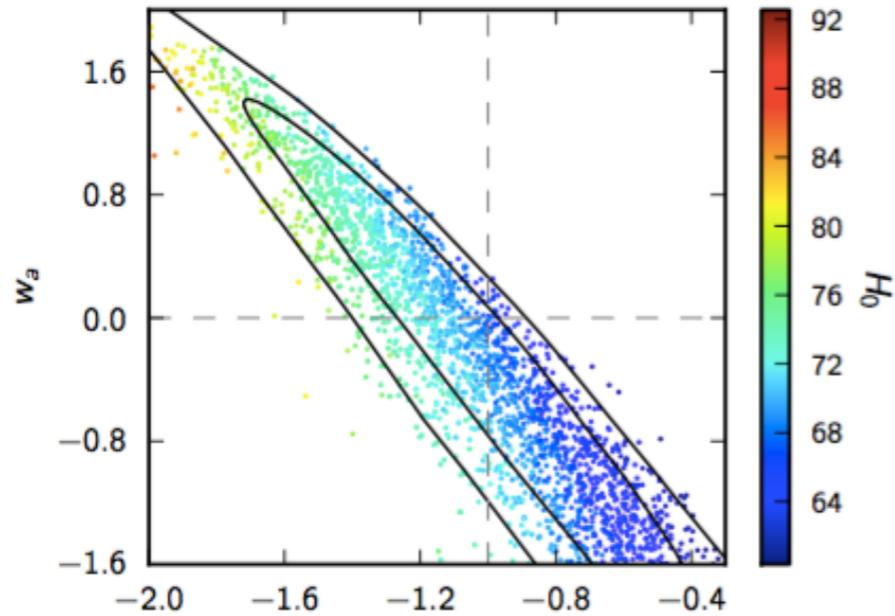


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# Dark Energy Nobel Prize 2011



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**Saul Perlmutter**



Photo: Belinda Pratten, Australian National University

**Brian P. Schmidt**



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**Adam G. Riess**

image from Nobelprize.org

in 1998 found that now

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$$\ddot{a} > 0 \quad \mathbf{!!!}$$

$$\frac{\ddot{a}}{a} = -\frac{4\pi G_N}{3} (\varepsilon + 3p)$$

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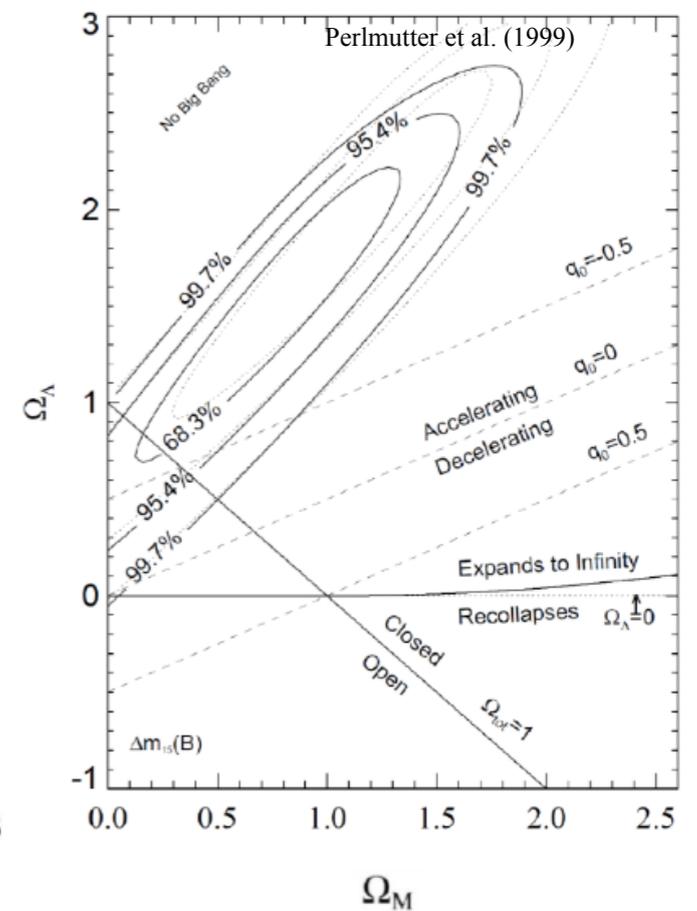
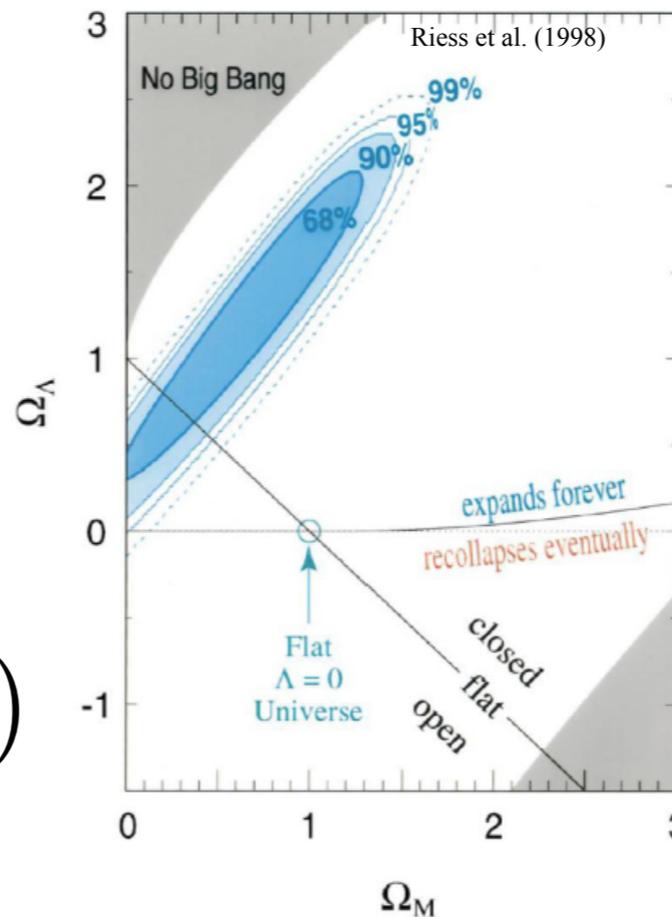
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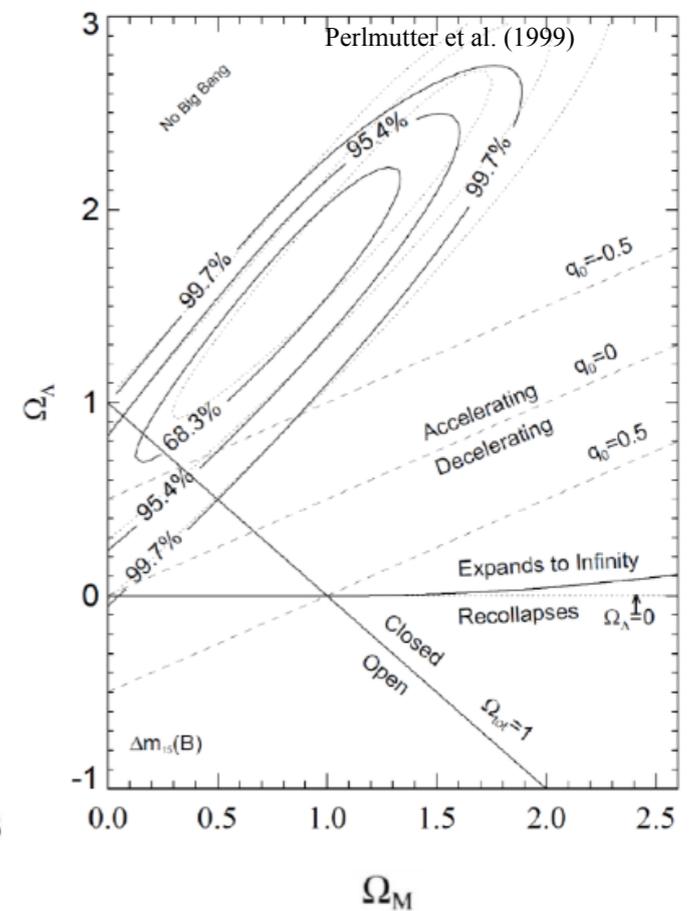
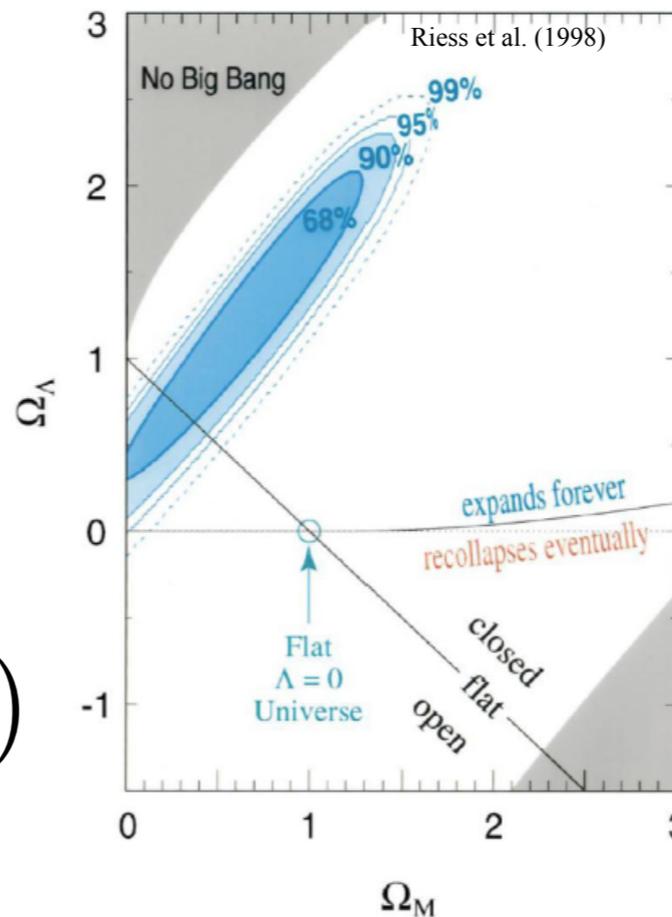
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# Negative pressure and Vacuum

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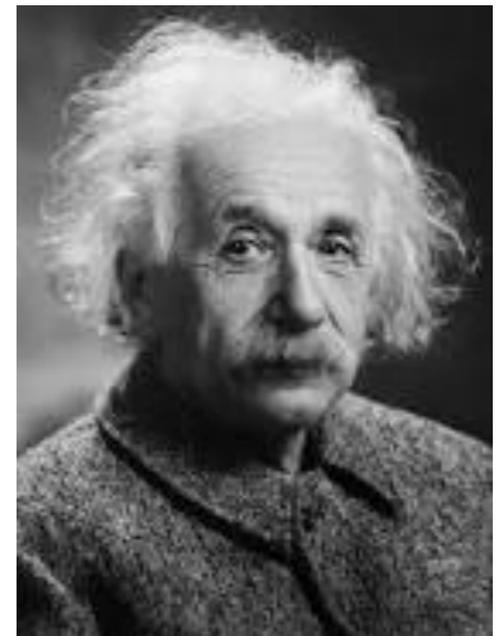
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Cosmological constant:  $\varepsilon = \Lambda$



(1917)

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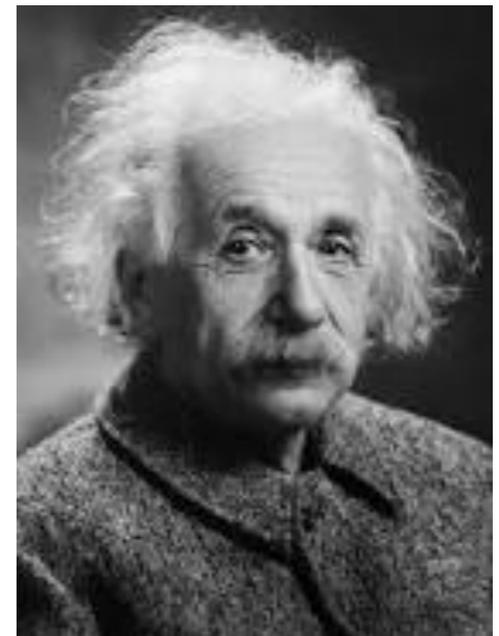
Euler relation

$$\varepsilon = Ts + \mu n - p \nearrow$$

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$$T_{\mu\nu} = \Lambda g_{\mu\nu}$$



(1917)

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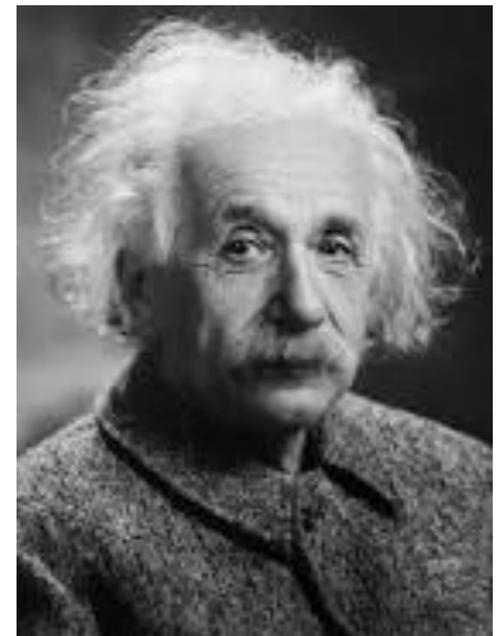
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(1917)

# Quantum Fluctuations

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Heisenberg uncertainty relation

$$\delta q \cdot \delta p \geq \frac{1}{2} \hbar$$



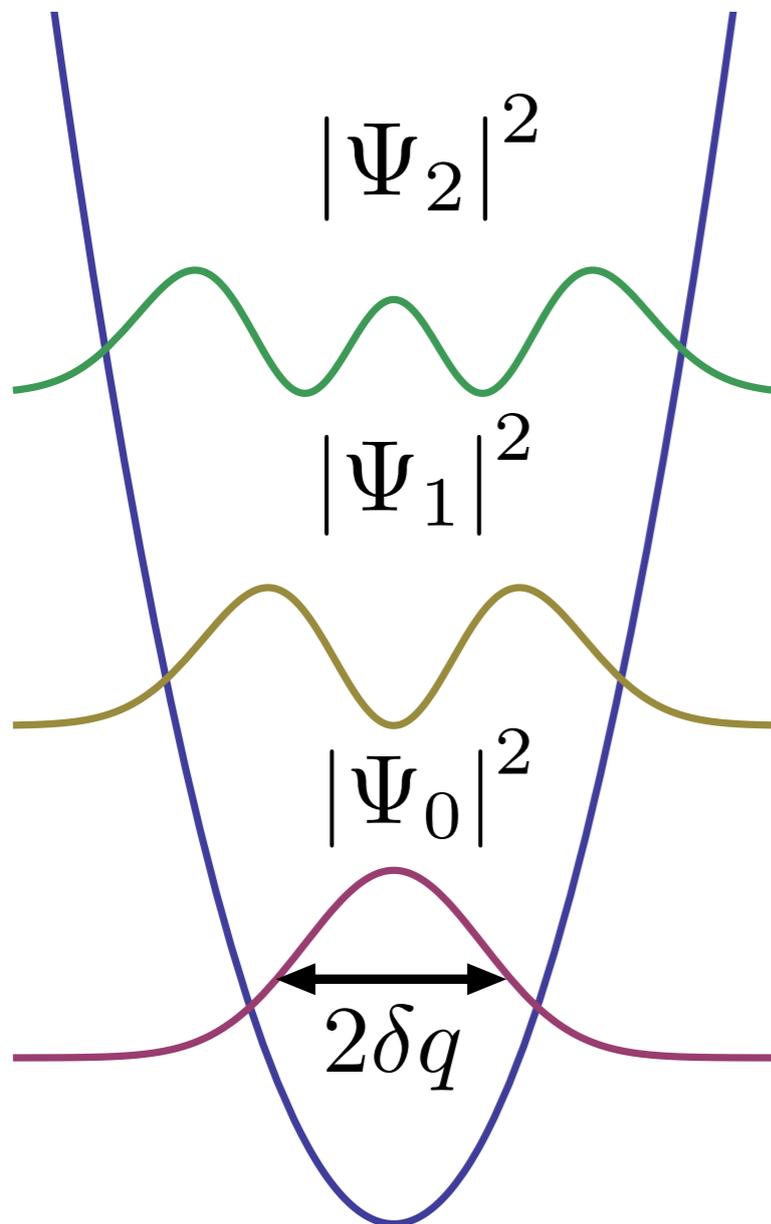
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$m=1$  Oscillator

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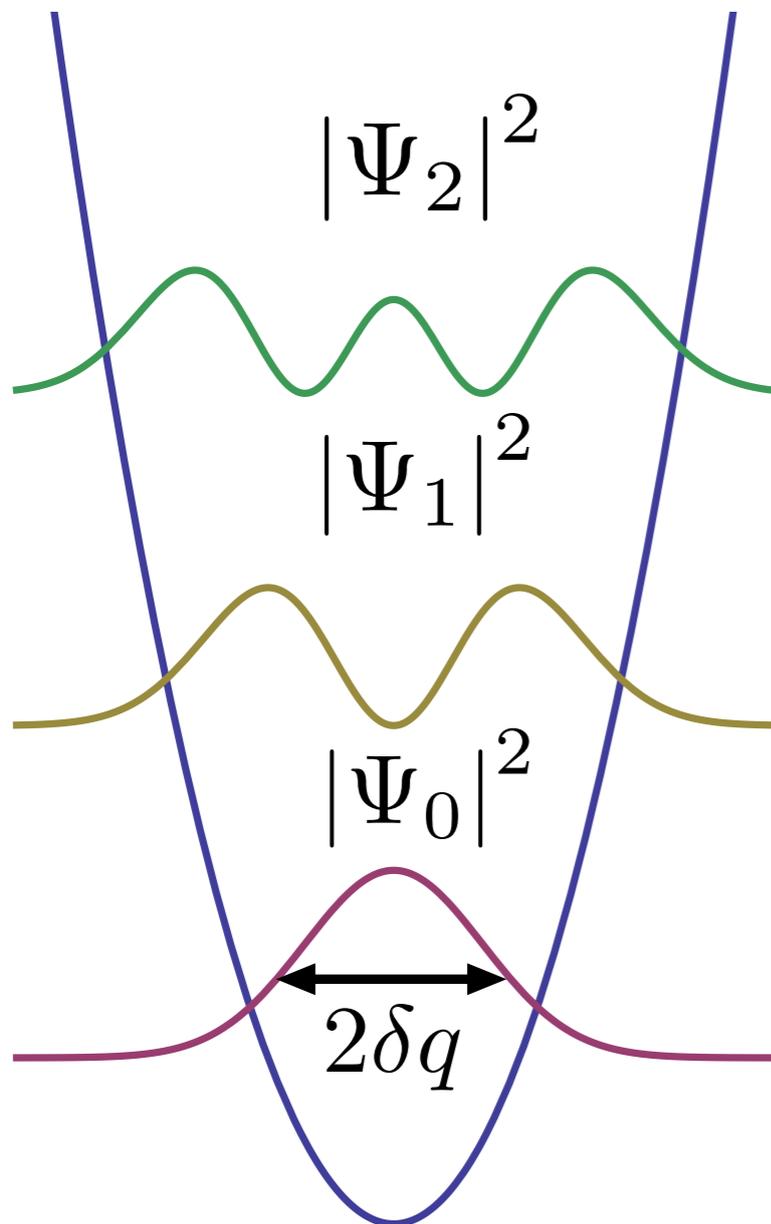
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vacuum also means  
*minimal possible* fluctuations!

$$\delta p = \omega \delta q \quad \rightarrow \quad \delta q = \sqrt{\frac{\hbar}{2\omega}}$$



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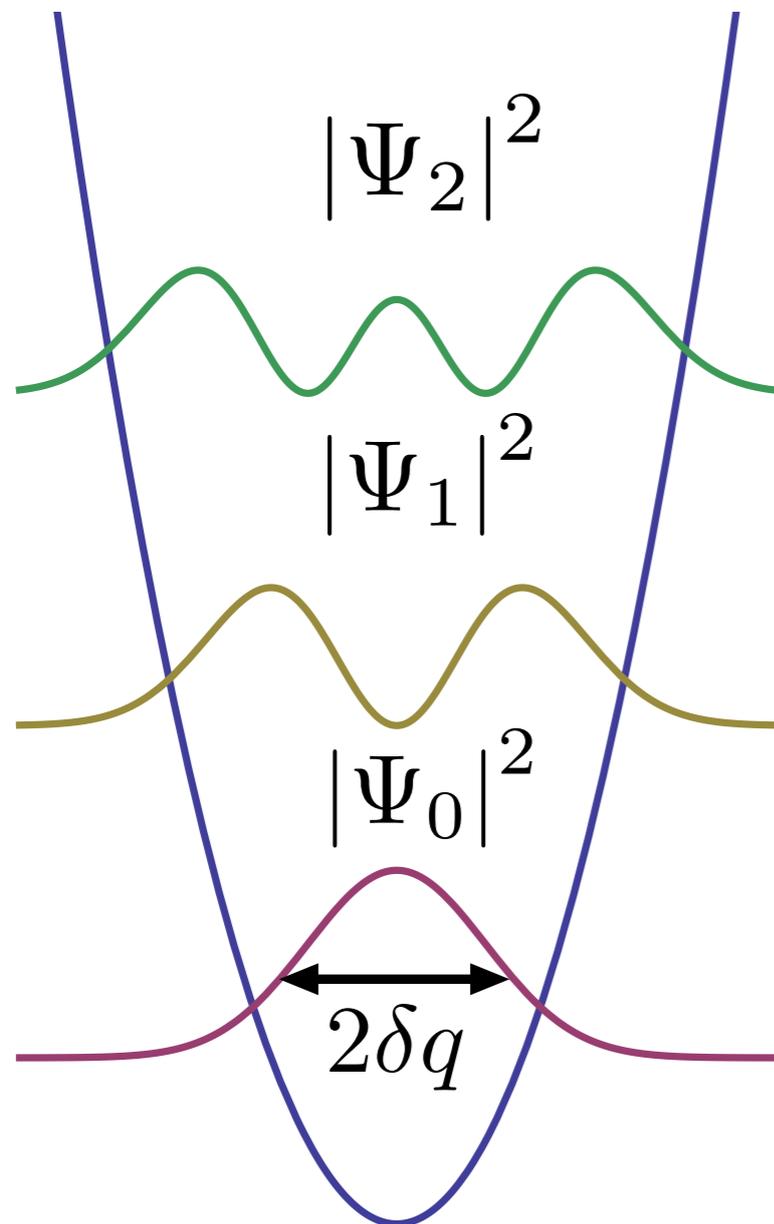


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$$\delta p = \omega \delta q$$

$$\delta q = \sqrt{\frac{\hbar}{2\omega}}$$

$$E_0 \simeq \frac{1}{2} (\delta p^2 + \omega^2 \delta q^2) = \omega^2 \delta q^2 = \frac{\hbar\omega}{2}$$

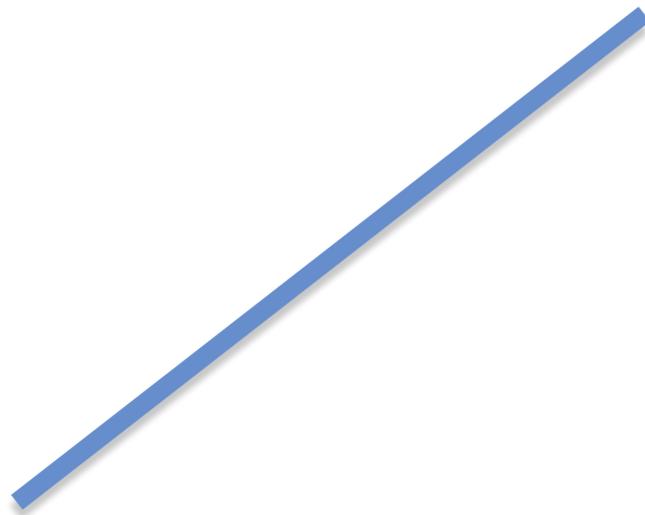
minimal possible “vacuum” energy

# Measuring Fields

- Every device has a *finite* resolution
- One cannot measure a field at a point in *space*
- Every measurement of a field yields a smoothed/  
*space-averaged / coarse grained* on some scale  $\ell$   
field  $\phi_\ell(x)$

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# Averaging fields

- A device with *finite adjustable* resolution  $\ell$  measures eigenvalues of the field operator

$$\hat{\phi}_\ell(\mathbf{x}, t) = \int d^d \mathbf{x}' W_\ell^\phi(\mathbf{x} - \mathbf{x}') \hat{\phi}(\mathbf{x}', t)$$

smearred by a family of window functions

$$W_\ell^\phi(\mathbf{x}) = \ell^{-d} \cdot w^\phi\left(\frac{\mathbf{x}}{\ell}\right)$$

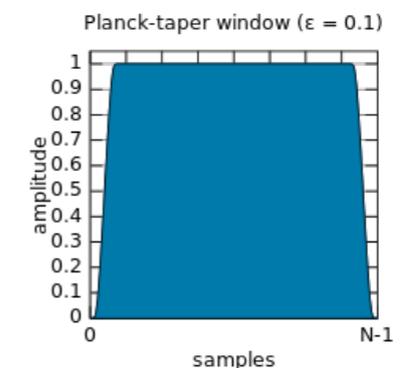
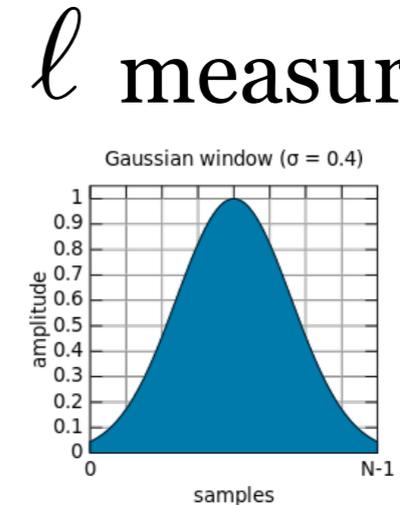
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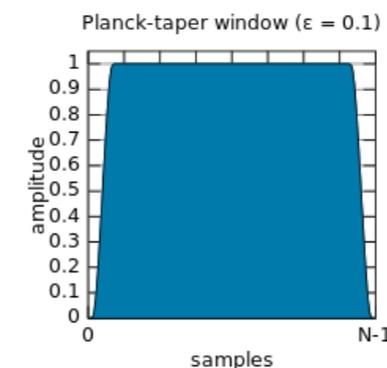
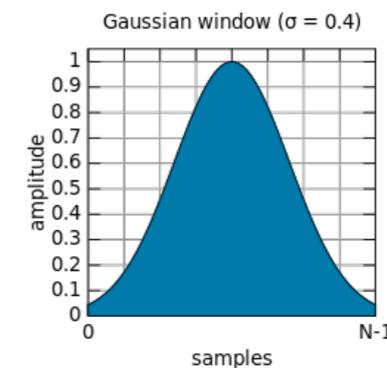
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$\ell$  measures



- Another device which we put *at the same place* measures canonical momentum  $\hat{p}(t, \mathbf{x})$  averaging it with some other family of window functions with the same scaling property above. Suppose that the resolution is the same.

# Canonical Quantisation

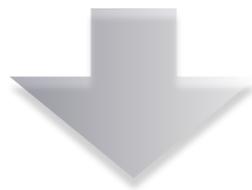
# Canonical Quantisation

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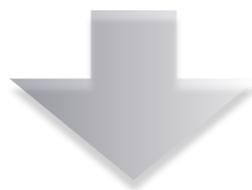
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$$\left[ \hat{\phi}_\ell(\mathbf{x}), \hat{p}_\ell(\mathbf{y}) \right] = i\hbar \cdot \ell^{-d} \cdot \mathcal{D} \left( \frac{\mathbf{x} - \mathbf{y}}{\ell} \right)$$

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where

$$\mathcal{D}(\mathbf{r}) = \int d^d \mathbf{r}' w^\phi(\mathbf{r} - \mathbf{r}') w^p(\mathbf{r}')$$

dimensionless

this convolution of shapes does not depend on scale  $\ell$  but only on the way of averaging

# Uncertainty Relations, QFT

fluctuations *on scale*  $\ell$ :

$$\delta_{\Psi} \Phi_{\ell} \equiv \sqrt{\langle \Psi | \hat{\Phi}_{\ell}^2 | \Psi \rangle - \langle \Psi | \hat{\Phi}_{\ell} | \Psi \rangle^2}$$

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$$\delta \phi_{\ell}(\mathbf{x}) \cdot \delta p_{\ell}(\mathbf{x}) \geq \frac{\hbar}{2} \cdot \mathcal{D}_0 \cdot \ell^{-d}$$

for two Gaussian window functions  $\mathcal{D}_0 = (2\sqrt{\pi})^{-d}$

$$\delta\phi_\ell(\mathbf{x}) \cdot \delta p_\ell(\mathbf{x}) \geq \frac{\hbar}{2} \cdot \mathcal{D}_0 \cdot \ell^{-d}$$



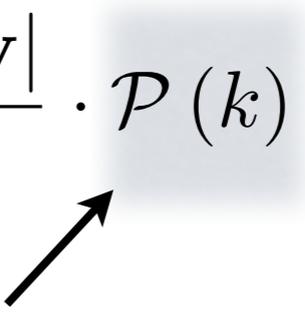
the shorter is the scale at which we look, the more quantum is the world

# Two -Point Function / Correlator and the Power Spectrum

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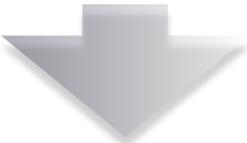
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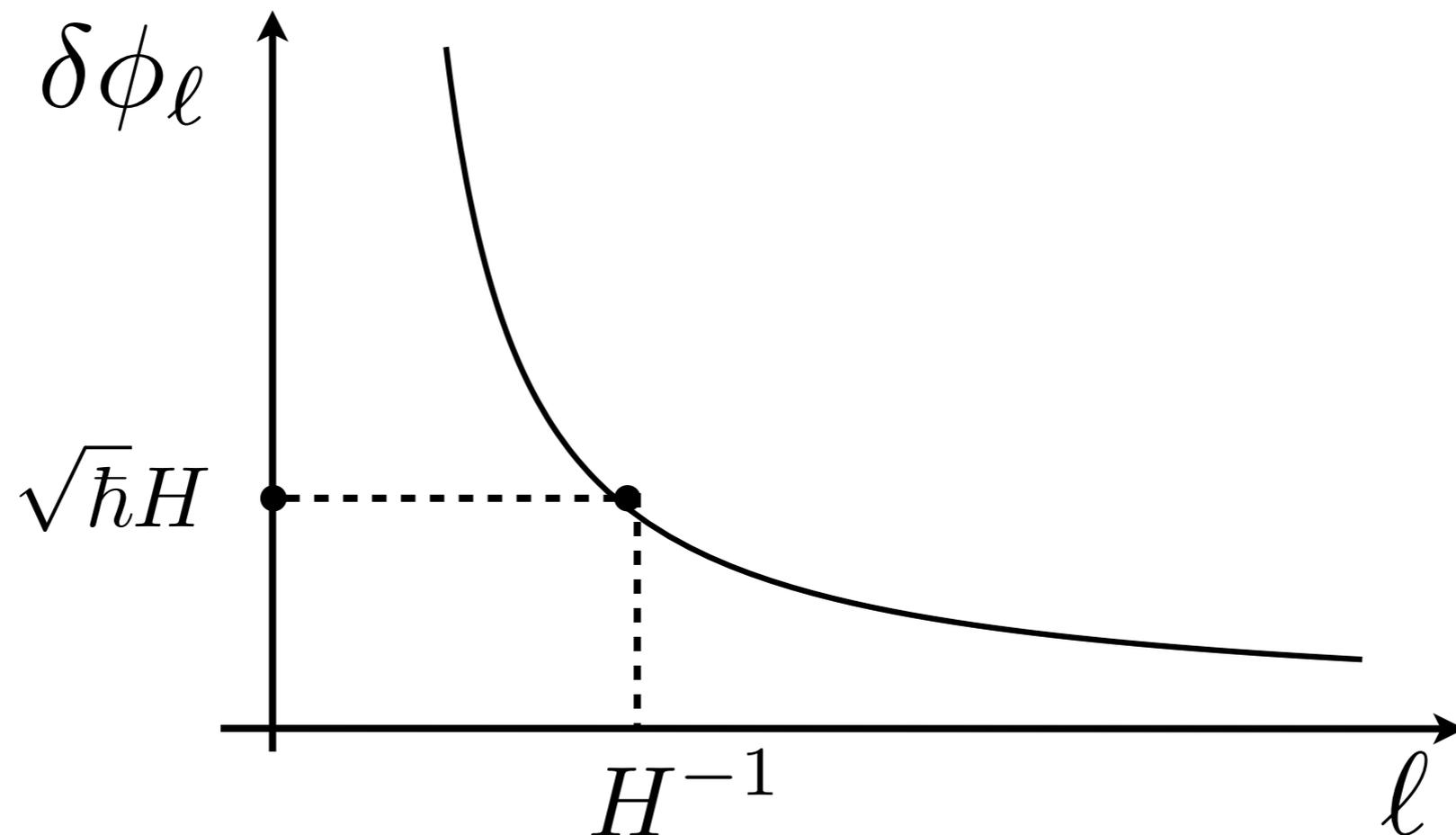
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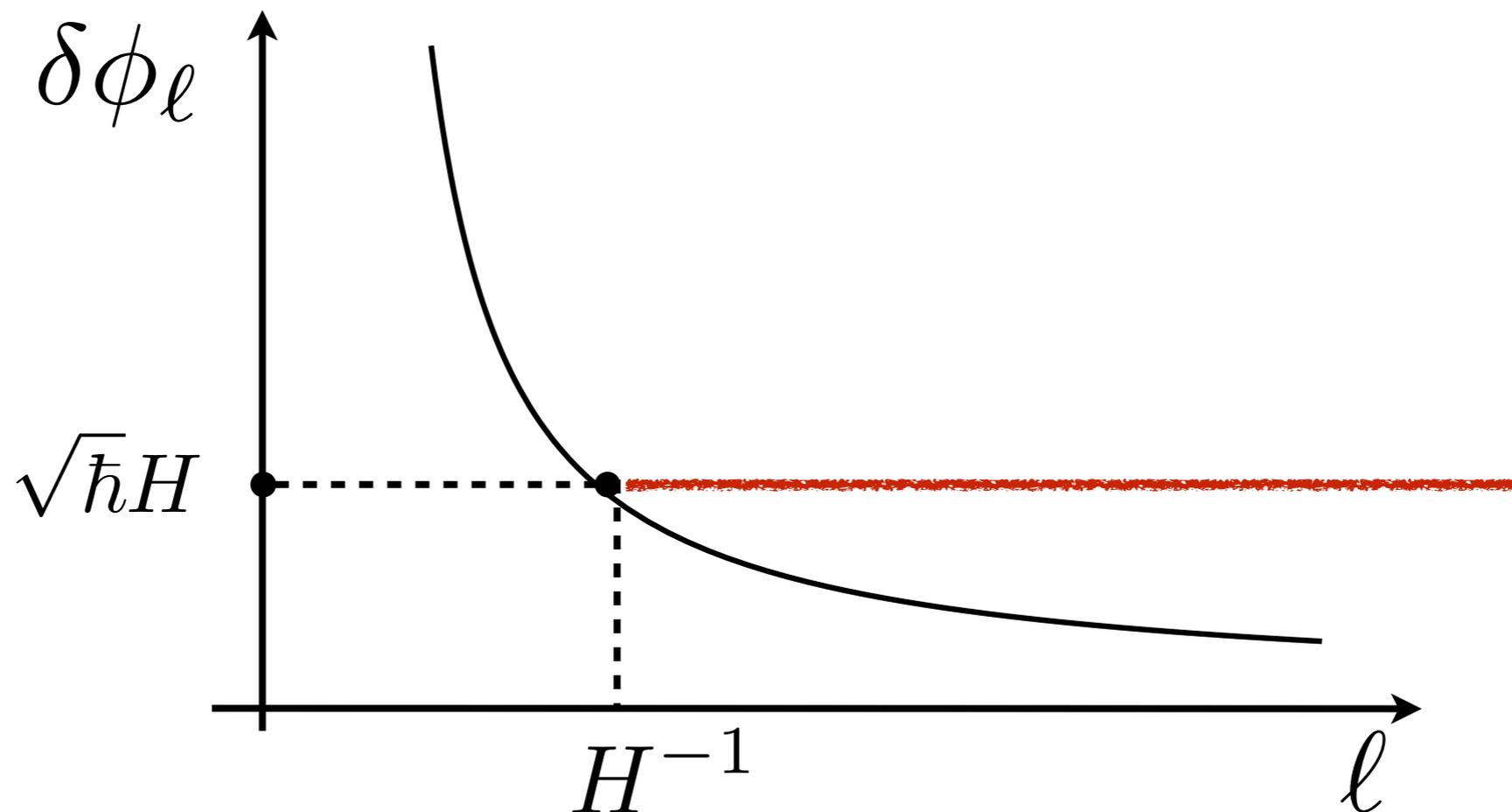
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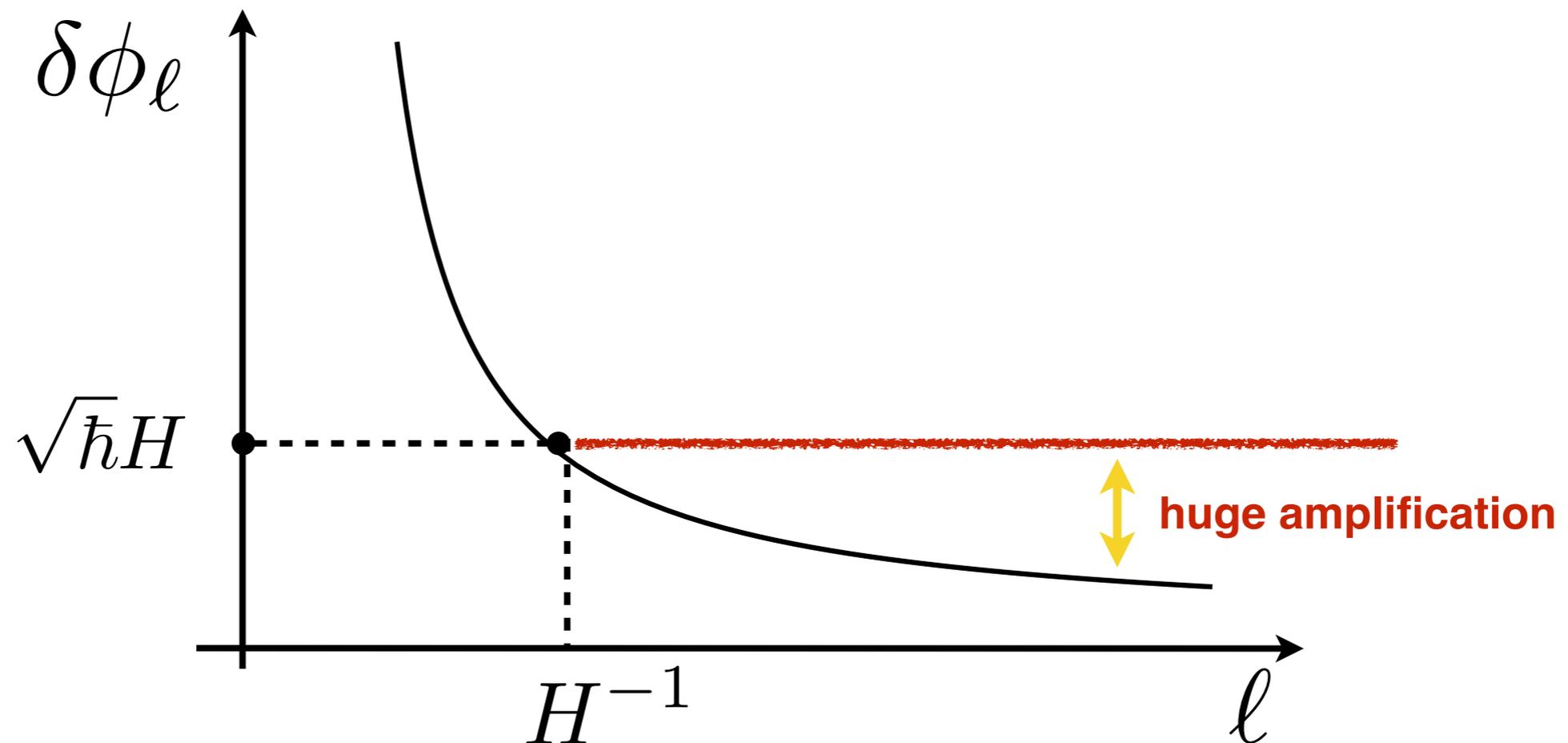
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For short range of scales one  
can parametrise  
the resulting Newtonian  
potential as a power law

$$\Phi_\ell \propto \ell^{(1-n_s)/2}$$

$$n_s < 1$$

spectral index,  
prediction slightly red

**Galaxies and  
Large Scale Structure  
do gravitate!**

**But all of them appeared out of  
quantum fluctuations!!!**



**quantum fluctuations should  
gravitate!!!**

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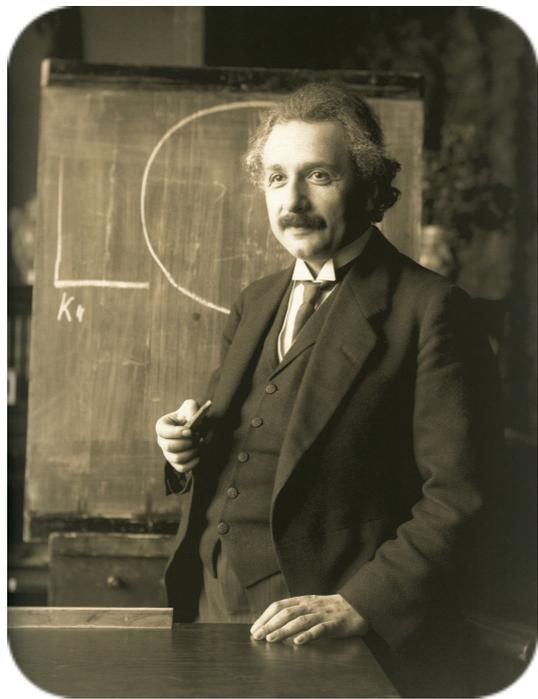
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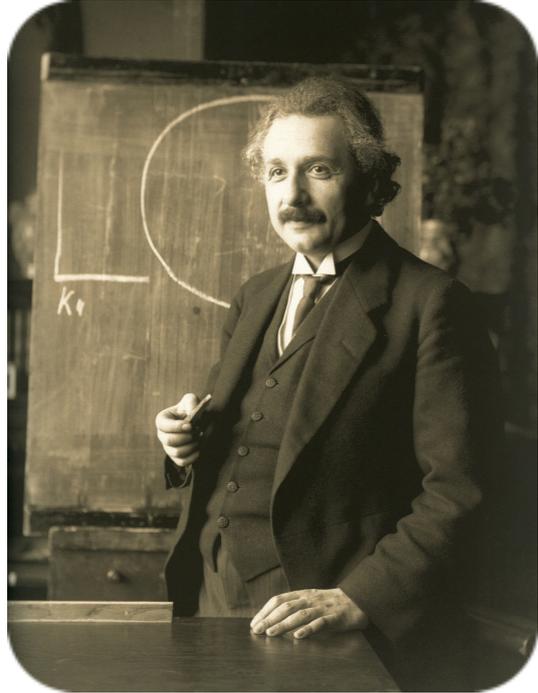
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EINSTEIN: Gravitationsfelder im Aufbau der materiellen Elementarteilchen 349  
 IN: Königlich Preussische Akademie der Wissenschaften (Berlin). Sitzungsberichte (1919): 349-356.

## Spielen Gravitationsfelder im Aufbau der materiellen Elementarteilchen eine wesentliche Rolle?

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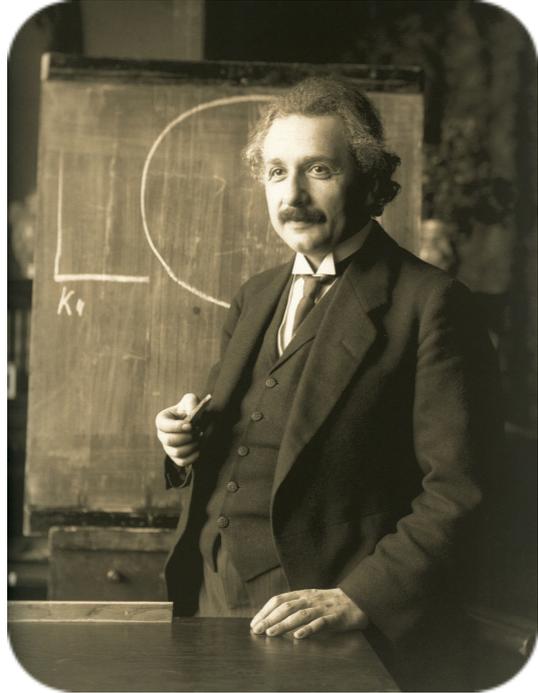
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neue universelle Konstante  $\lambda$  eingeführt werden mußte, die zu der Gesamtmasse der Welt (bzw. zu der Gleichgewichtsdichte der Materie) in fester Beziehung steht. Hierin liegt ein besonders schwerwiegender Schönheitsfehler der Theorie.

$$G_{\mu\nu} - T_{\mu\nu} - \Lambda g_{\mu\nu} = 0$$

# Decoupling vacuum energy from spacetime curvature

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invariant under vacuum shifts of  
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What is the action for  
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